

# The force hierarchy and geodynamo regimes of Earth's core



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Thanks to collaborators:

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University  
of Glasgow



Second oldest university in Japan (1897)



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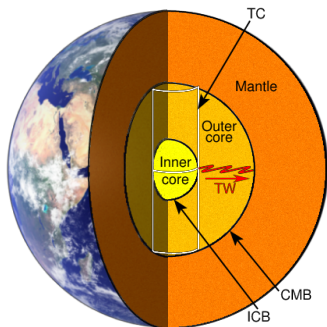
Interest in the dynamics of Earth's core arises from its ability to generate the geomagnetic field.

What might you already know about the geomagnetic field?

Field is predominantly dipolar (but also note the patches of reversed flux found at high latitudes).

Reversals of the field dipolarity occur (seemingly at random intervals and a reversal takes thousands of years to complete).

# Structure of Earth



ICB = Inner core boundary  
 CMB = Core-mantle boundary  
 TC = Tangent cylinder

Fluid outer core is seat of dynamo giving rise to geomagnetic field.

Convection arises from heat and light material released at inner core boundary.

Magnetic field is continually replenished through induction (combining Faraday's law, Ampere's law, and Ohm's law)

Twisting and stretching of field lines by chaotic convection generates electric current, in turn re-generating magnetic field.



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Aim to match simulations to observations of the changing geomagnetic field thereby understanding dynamics in the core.

# Geodynamo simulations - physical setup

Spherical polar coordinate system  $(r, \theta, \phi)$ .

Spherical shell radially bounded above at  $r = r_o$  by an electrically insulating mantle and below at  $r = r_i$  by an electrically insulating (or conducting) inner core.

Rotates about the vertical  $z$ -axis with rotation rate  $\Omega$  and gravity acts radially inward,  $g = -g\hat{r}$ .

Boussinesq approximation used - density,  $\rho$ , treated as a constant except for the source of buoyancy

Fluid is assumed to have constant values of  $\nu$ ,  $\kappa$ , and  $\eta$ , the outer core density, kinematic viscosity, thermal diffusivity and magnetic diffusivity respectively.

# Geodynamo equations

Evolution equations for velocity,  $u$ , temperature  $T$ , and magnetic field,  $B$ :

$$\underbrace{E_m \frac{\partial u}{\partial t}}_{\text{inertia (I)}} + \underbrace{u \cdot \nabla u}_{\text{inertia (I)}} = \underbrace{-\nabla p}_{\text{pressure (P)}} + \underbrace{2\mathbf{\Omega} \times u}_{\text{Coriolis (C)}} + \underbrace{(\nabla \times \mathbf{B}) \times \mathbf{B}}_{\text{Lorentz (M)}} + \underbrace{\frac{\rho_a T}{\rho}}_{\text{buoyancy (A)}} + \underbrace{E \nabla^2 u}_{\text{viscous (V)}};$$

$$\frac{\partial T}{\partial t} + u \cdot \nabla T = \alpha \nabla^2 T;$$

$$\frac{\partial \mathbf{B}}{\partial t} = \nabla \times (u \times \mathbf{B}) + \eta \nabla^2 \mathbf{B};$$

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$$\frac{\partial T}{\partial t} + u \cdot \nabla T = \alpha r^2 T;$$

$$\frac{\partial \mathbf{B}}{\partial t} = \nabla \times (\mathbf{u} \times \mathbf{B}) + \alpha r^2 \mathbf{B};$$

with conditions:  $r \cdot u = 0$  and  $r \cdot B = 0$ .

4 key input parameters:

$$E = \frac{1}{d^2}; \quad [E]_{\text{core}} = 10^{15}; \quad [E]_{\text{sim}} = 2 [10^6; 10^3]$$

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4 key input parameters:

$$E = \frac{g \alpha D^3}{\nu \kappa}; \quad [E]_{\text{core}} \approx 10^{15}; \quad [E]_{\text{sim}} \approx [10^6; 10^3]$$

$$E_m = \frac{g \alpha D^3}{\nu}; \quad [E_m]_{\text{core}} \approx 10^9; \quad [E_m]_{\text{sim}} \approx [10^7; 10^3]$$

$$Ra = \frac{g \alpha D^3}{\nu \kappa}; \quad [Ra]_{\text{core}} \approx 10^{10}; \quad [Ra]_{\text{sim}} \approx [10^8; 10^4]$$

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$$\frac{\partial T}{\partial t} + u \cdot \nabla T = q r^2 T;$$

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with conditions:  $r \cdot u = 0$  and  $r \cdot B = 0$ .

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$$\kappa_a = \frac{g}{d}; \quad [\kappa_a]_{\text{core}} = 10^{10}; \quad [\kappa_a]_{\text{sim 2}} = [10^{-10}; 10^4]$$

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$$\alpha = \frac{g \Delta T d}{\kappa}; \quad [\alpha]_{\text{core}} = 10^{10}; \quad [\alpha]_{\text{sim}} = 2 \times [10^{-1}; 10^4]$$

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$Ra = \frac{g \Delta T d}{\kappa} = \alpha = q$  is an alternative Rayleigh number (useful to relate to non-magnetic problem).

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$Ra = \frac{g \Delta T d}{\nu \kappa} = Ra_m = q Ra_c$  is an alternative Rayleigh number (useful to relate to non-magnetic problem).

Often use  $Ra^0 = Ra = Ra_c = Ra_m = q Ra_c$ , as a measure of supercriticality.  $Ra_c$  is the critical Rayleigh number for the onset of (non-magnetic) convection.

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  - ! Can have unintended outcomes such as non-dipolar solutions, solutions with a weak magnetic field, etc.
- 2 Correct solution space: aim to find solutions with the expected balance of forces within the momentum equation by performing parameter sweeps
  - ! Allows for the identification of suitable parameter regimes despite input parameters not close to Earth-like values.  
Then preserve the force balance by moving all parameters towards Earth-like values in a systematic way.

# Force balances

Forces acting in the non-dimensionalised system are:

$$F_I = E_m \frac{\partial u}{\partial t} + u \cdot r \cdot u$$

$$F_P = r \cdot p$$

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Geostrophic balance  $F_P = F_C$ . For no magnetic field and no convection, (and small  $E$  and  $E_m$ ).  
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VAC balance,  $F_V \sim F_A \sim F_C$ . Close to onset of convection/dynamo action - actually preserves geostrophy due to form of  $V$  and  $A$ .

CIA balance,  $F_C \sim F_I \sim F_A$ . At high convective driving and weak magnetic field.



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Some previous investigations of force balances:

Rotvig & Jones, Phys Rev E, 2002

Soderlund+, PEPS, 2015

Yadav+, PNAS, 2016

Schaefer+, GJI, 2017

Schwaiger+, GJI, 2019, 2021

Teed & Dormy, JFM, 2023

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But that is not the case!

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Teed & Dormy (2023) proposed forming 'solenoidal forces' by directly eliminating gradient parts of all forces to observe the important 1st order balance (MAC, VAC, etc.) directly

$F = r \cdot A + r \cdot \nabla$ ; eliminate  $r \cdot \nabla$  by:

curling  $F$ . (Note: Taylor-Proudman constraint is formed this way!)

projecting of forces onto their solenoidal part:  $r \cdot A$

! Perform curl  
Remove gradient parts

Teed & Dormy, 2023

## Weak and strong field branches - theory

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Potential bistability - different regimes at the same input parameters

# Weak and strong eddy branches - simulations

Dormy, 2016; identification of strong eddy branch and bistability in DNS at  $E = 3 \times 10^4$  and  $E_m = 1.7 \times 10^5$

Requires  $E_m$  to be chosen within a 'sweet-spot' range of values (dependent on  $E$ )...

# Regime diagrams

Each plot is decreasing  $E_g$

Decreasing  $E_m$

Solutions become  
non-dipolar

Solutions become  
non-magnetic

Solutions become  
non-dipolar with increasing  $Ra^0$

# Bifurcation diagrams

Bifurcation diagrams for  
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Dormy, 2025

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Diagrams differ (isola, supercritical, subcritical bifurcations) as  $E_m$  is varied

Weak-strong branching is found at low enough  $E_m$

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Dormy, 2025

Tentative 3D bifurcation diagram for  $x_{edE}$

Dormy, 2025



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Tentative 3D bifurcation diagram for  $x \in \mathbb{E}$

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Supercritical branch exhibits a sharp step announcing the cusp

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Need to study dependence of such solutions on  $E_s$  and  $E_m$  to help determine constants

## Some questions to address

Now we'll look at some results on (solenoidal) forces (Teed & Dormy, 2023) and upcoming work on geodynamo branches as  $\text{Eis}$  lowered (Teed & Dormy, 2025).

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- 3 How does the branching between **weak and strong regimes** persist/scale as parameters are moved towards Earth-like values? I.e. lower  $\epsilon$  and  $E_m$ .

# Branches of dynamo action $E (= 10^4)$

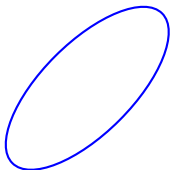
# Branches of dynamo action $E (= 10^4)$

At higher  $E_m$ :

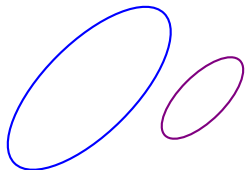


Branches of dynamo action  $(E_m = 10^4)$ 

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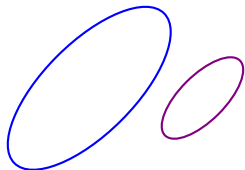


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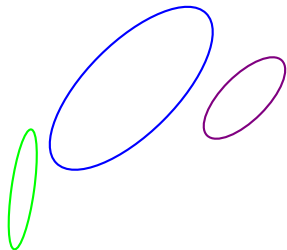
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# Branches of dynamo action $E_m (= 10^4)$



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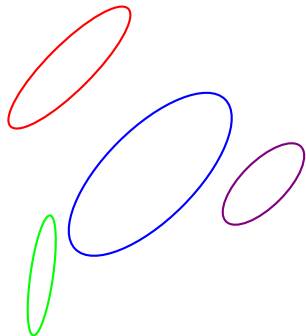
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Weak field dipolar regime  
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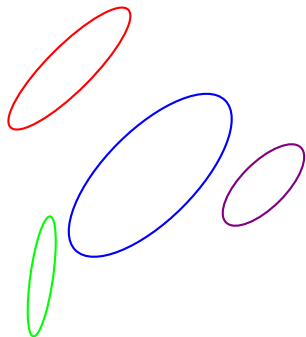
At lower  $E_m$ :

Weak eld dipolar regime

( $0 \ 1$ )...

...transitions to strong eld dipolar regime ( $0 \ 1$ )

# Branches of dynamo action $E_m (= 10^4)$



At higher  $E_m$ :

Dipolar 'Strongish' dipolar regime...  
 ...transitions to multipolar regime at  
 large enough  $Ra$

At lower  $E_m$ :

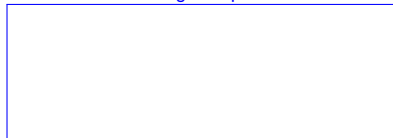
Weak old dipolar regime  
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...transitions to strong old  
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Bistability between weak and  
 strong branches in a region of  
 Ra-space

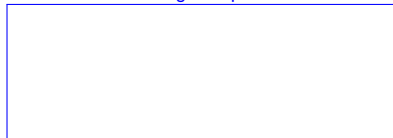
Typical regimes ( $aE = 10^{-4}$ )

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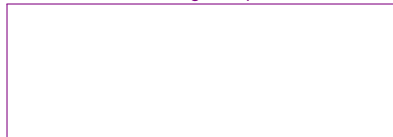


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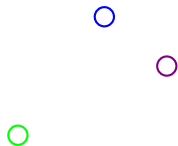
Strongish dipolar



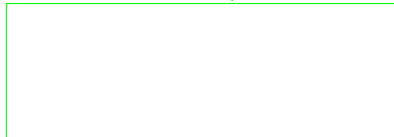
Fluctuating multipolar





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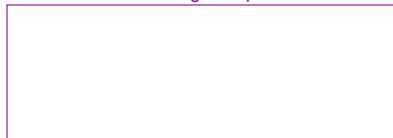
Weak eld dipolar

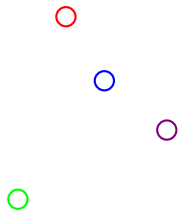


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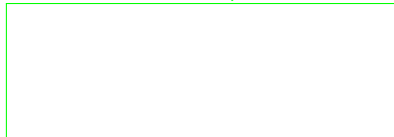


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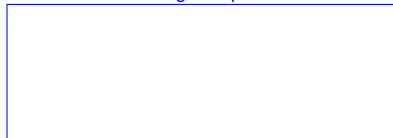


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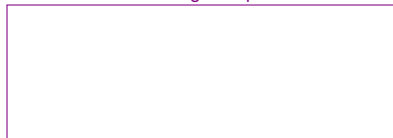
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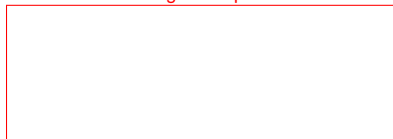
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# Comparing forces and solenoidal forces

! Perform curl  
Remove gradient parts

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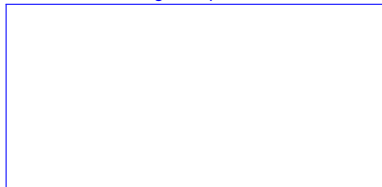


Forces show mostly QG balance with  $F_C^{ag} > F_C$  at some scales. Strange!

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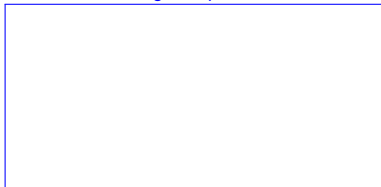
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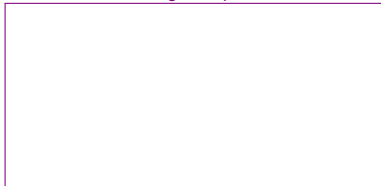
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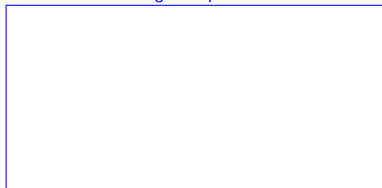


Forces show inertia entering zeroth order balance

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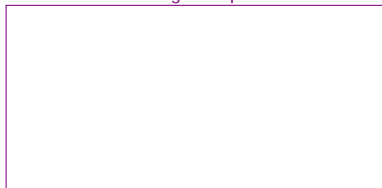
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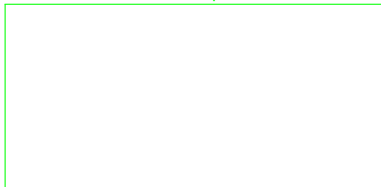
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Solenoidal forces reveal clear leading order CIA balance for multipolar regime

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Weak field dipolar



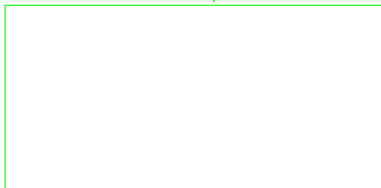
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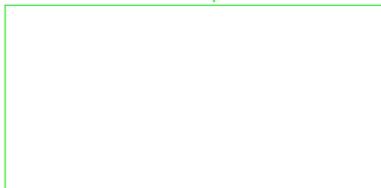
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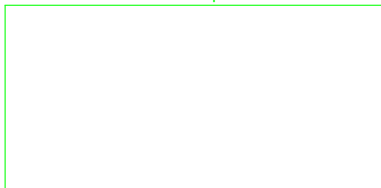


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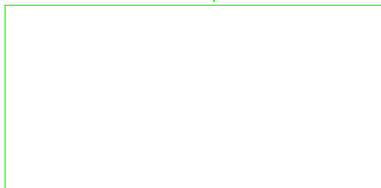
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Solenoidal forces reveal clear leading order MAC balance at large scale for strong field regime

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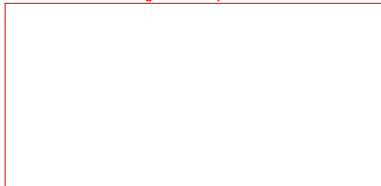
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Usefulness of ageostrophic Coriolis force lost at lengthscales where balance is not QG

# Solenoidal forces

*Strongish*

*Multipolar*

*Weak  $\ell$*

*Strong  $\ell$*