



Homogeneous nucleation experiment of cosmic dust

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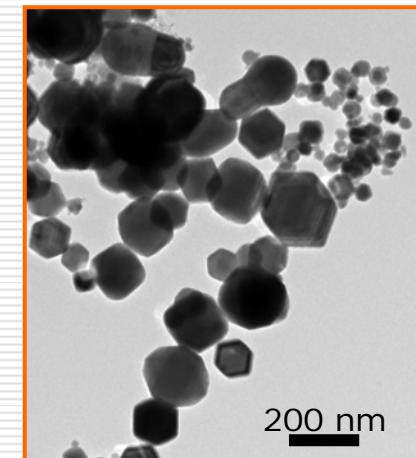
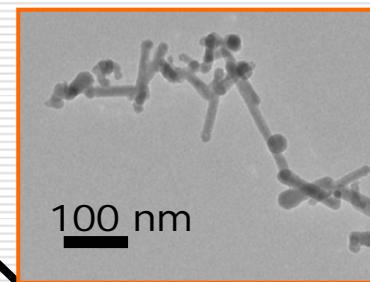
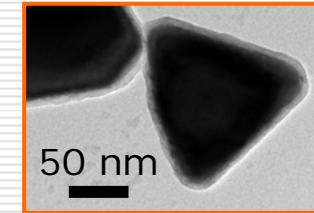
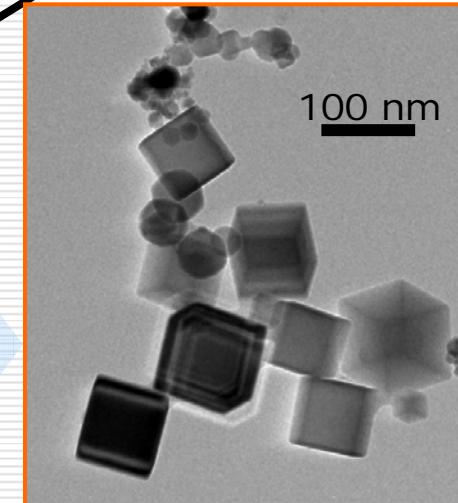
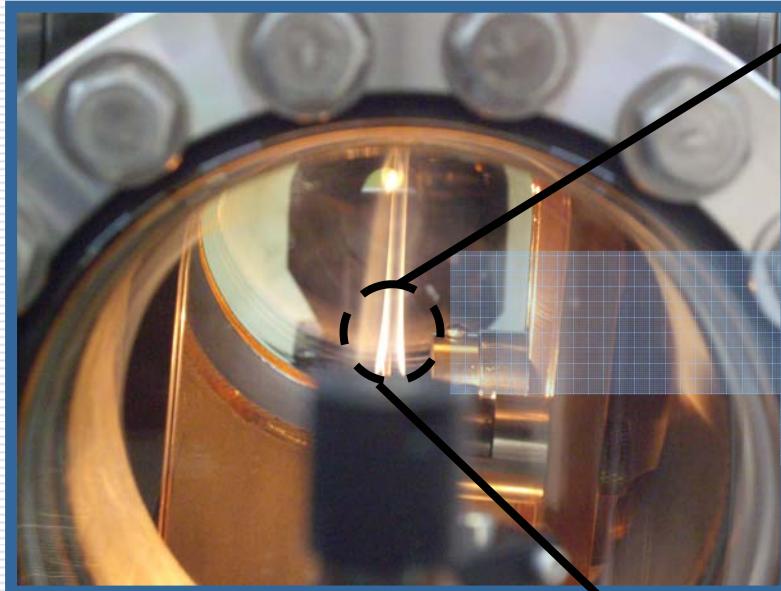
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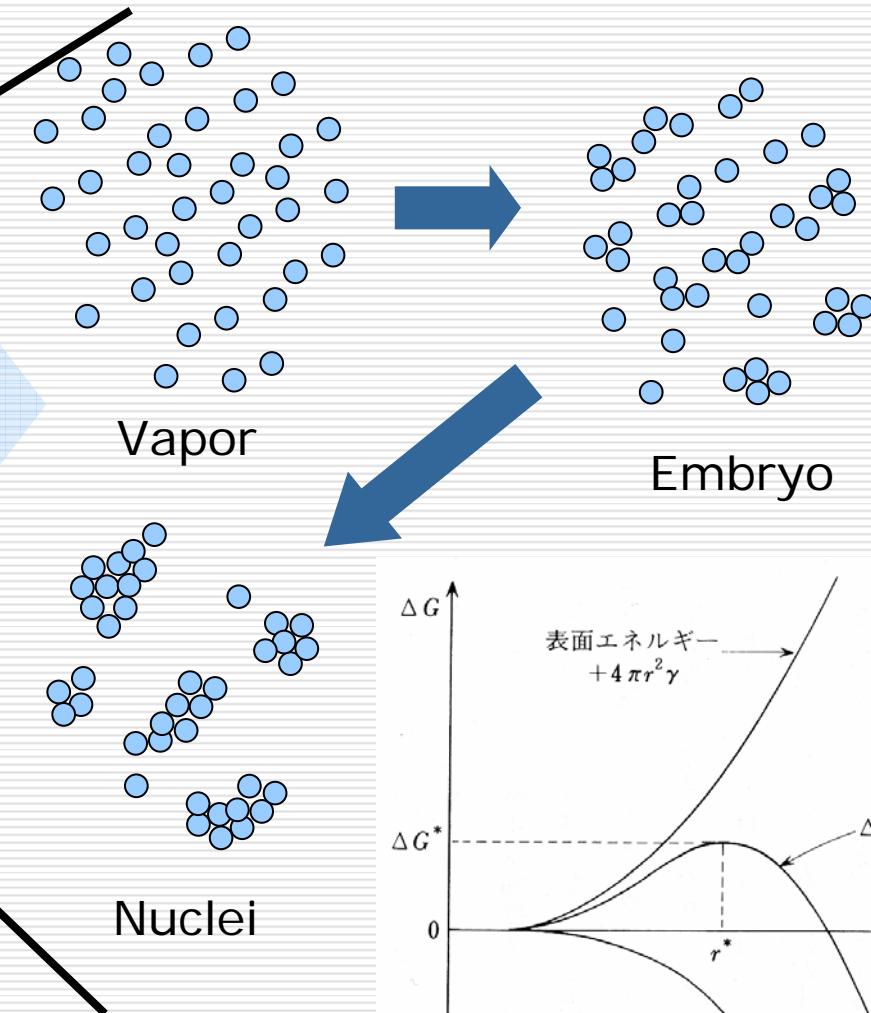
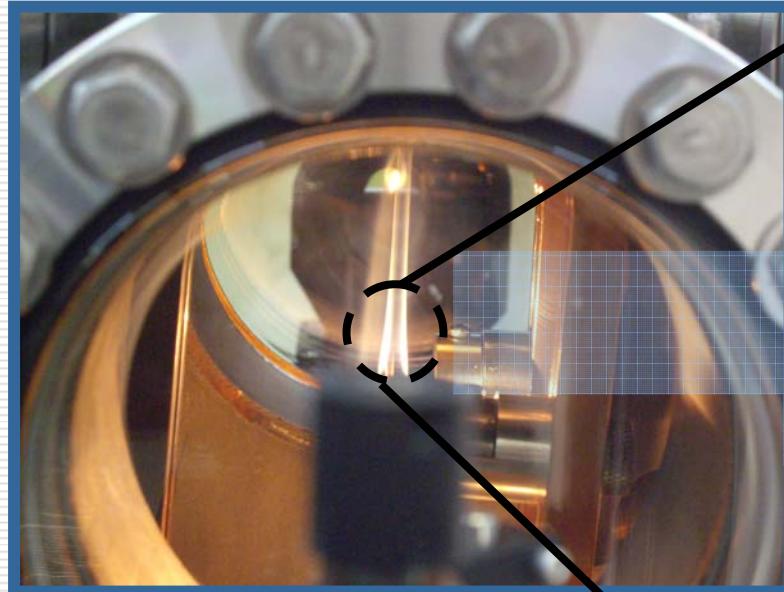
Nucleation in Lab & Space



Why nucleation?

- Number
- Morphology
- Crystal habit
- Size
- Size distribution
- Polymorphism

Nucleation in Lab & Space



Why nucleation?

- Number
- Morphology
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黒田登志雄
結晶は生きている

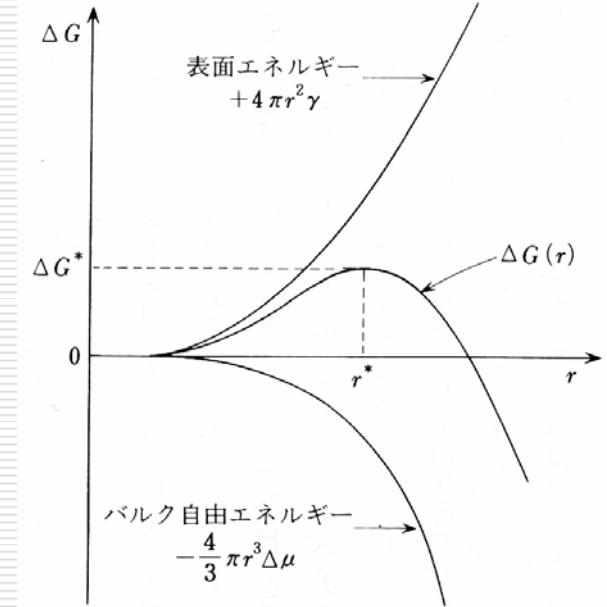
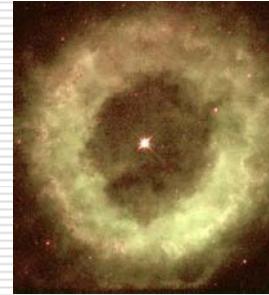
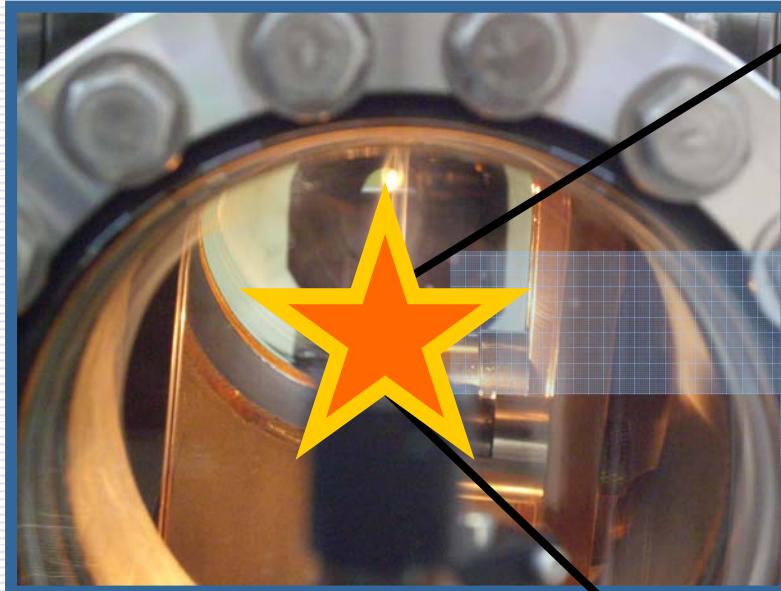
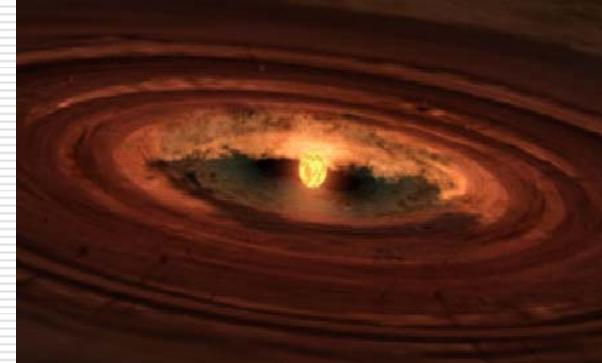


図 2.4 臨界核の半径 r^* と核生成に必要な自由エネルギー変化 ΔG^* .

Nucleation in Lab & Space



Evolved Star

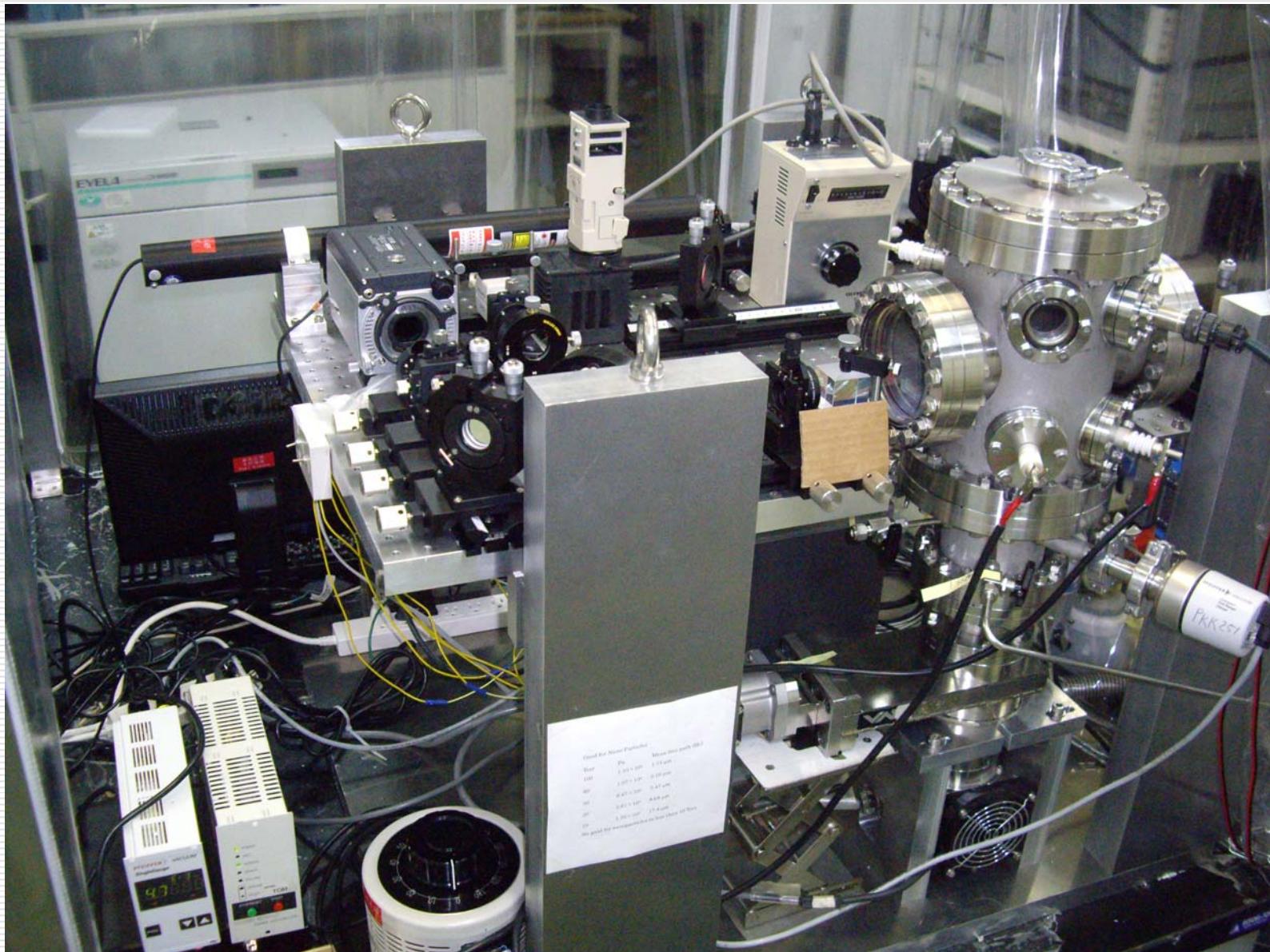


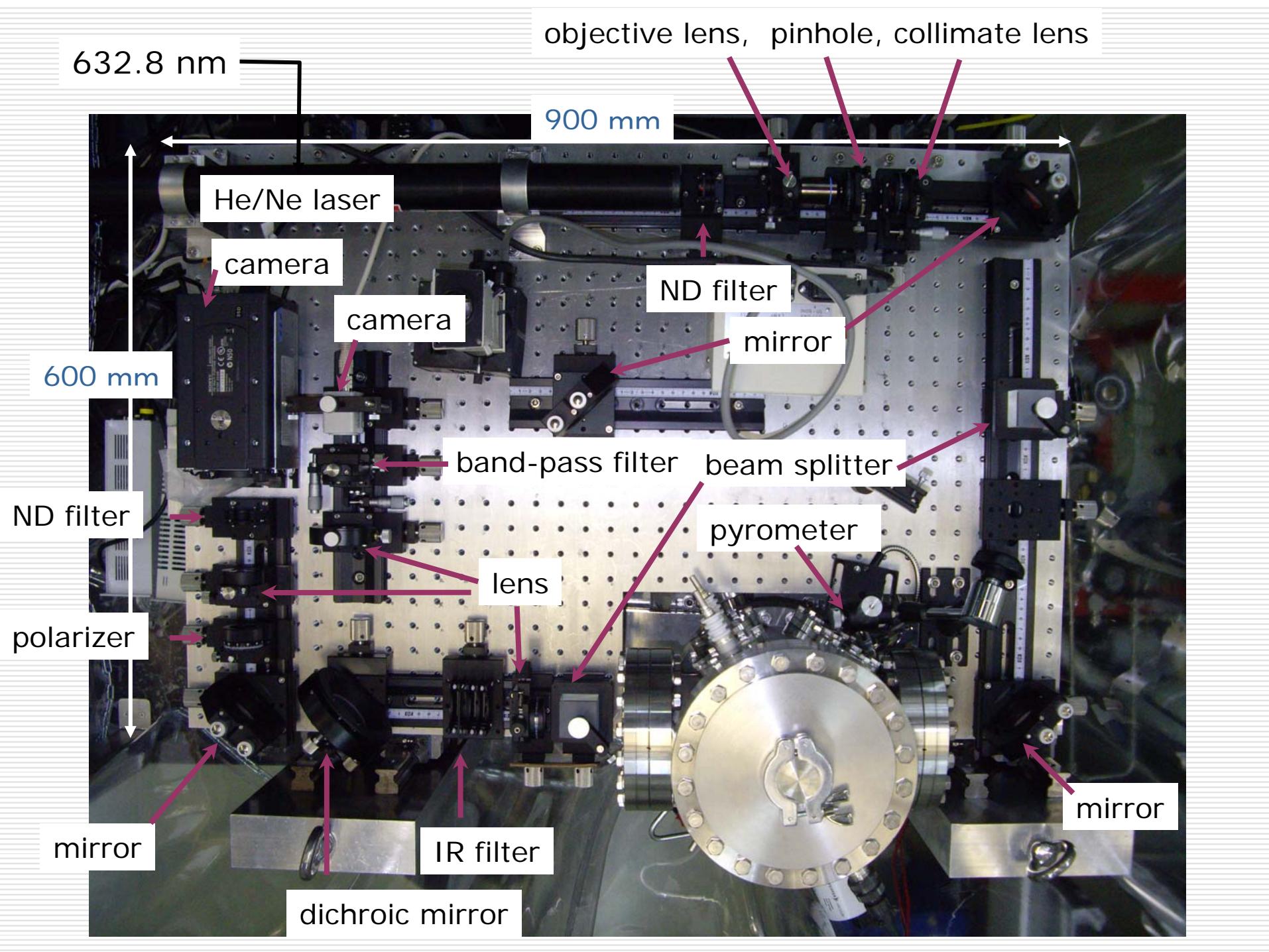
Planetary Nebula

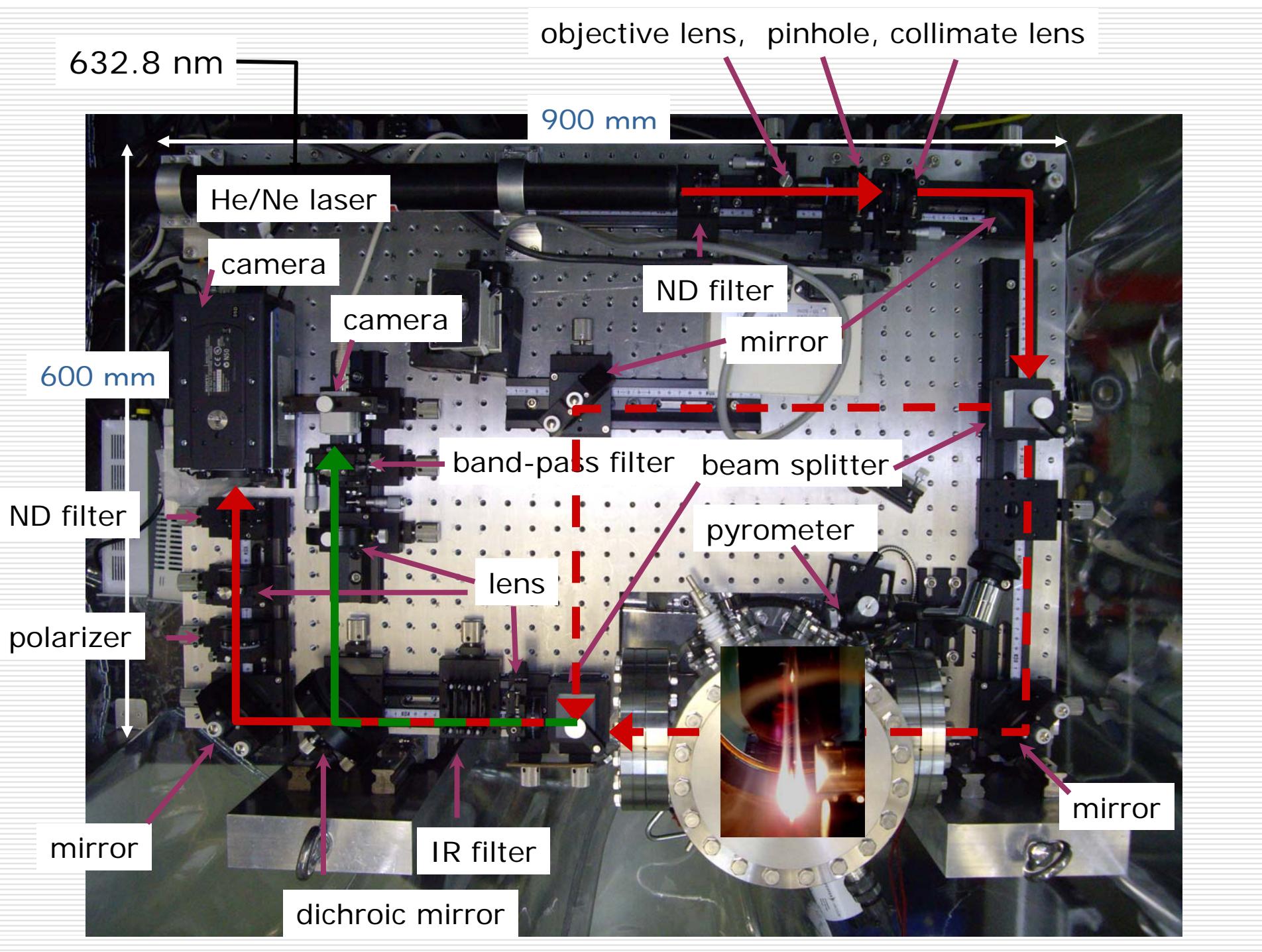
Why nucleation?

- Number
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- Crystal habit
- Size
- Size distribution
- Polymorphism

Chamber and optics

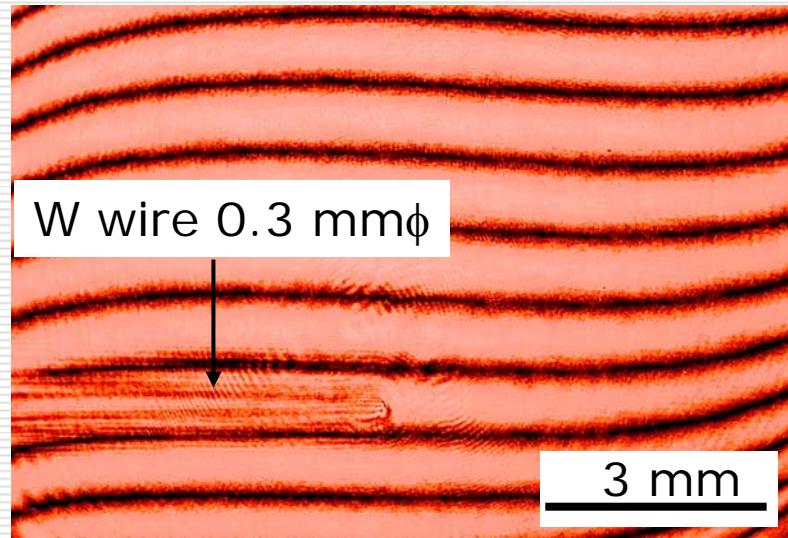




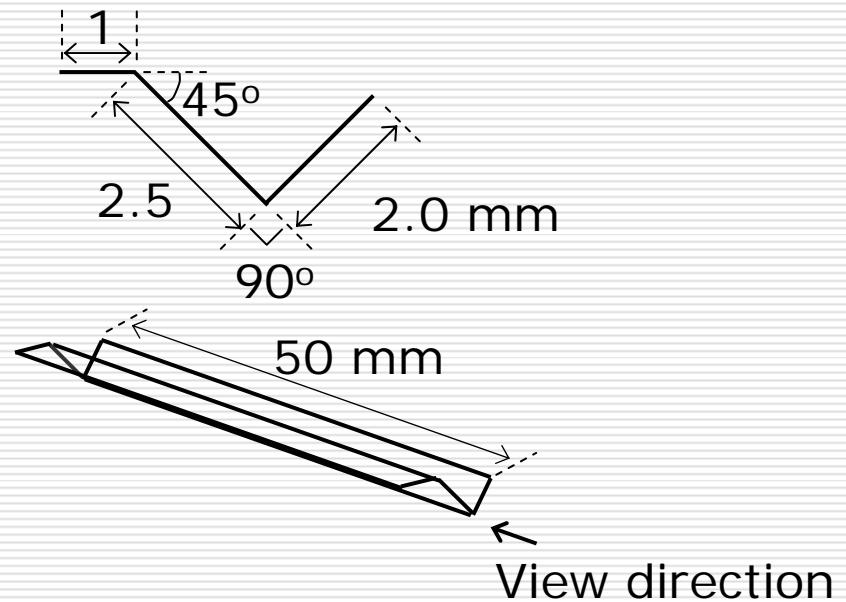
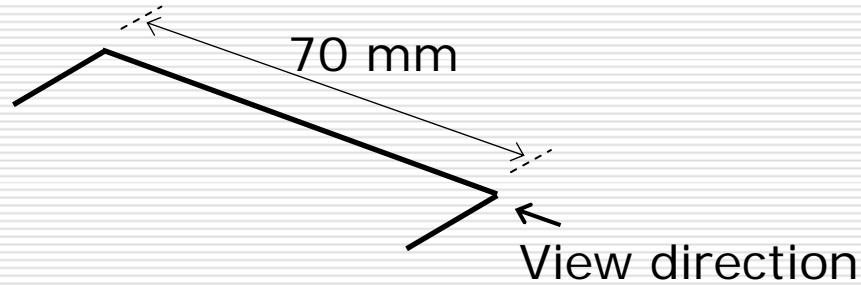
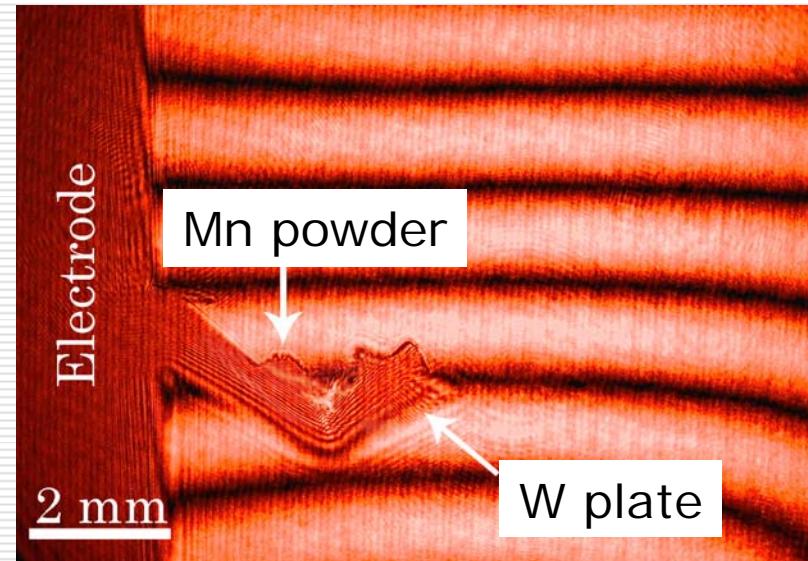


Interferogram

WO₃: Ar 9×10^3 , O₂ 1×10^3 Pa



Mn: Ar 2×10^4 Pa



Interferogram

Temperature: 298 K (25°C)

Gas: Ar 1×10^4 Pa

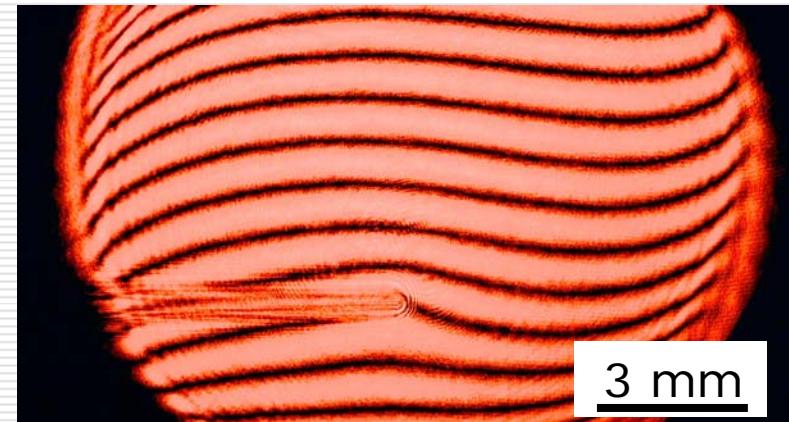
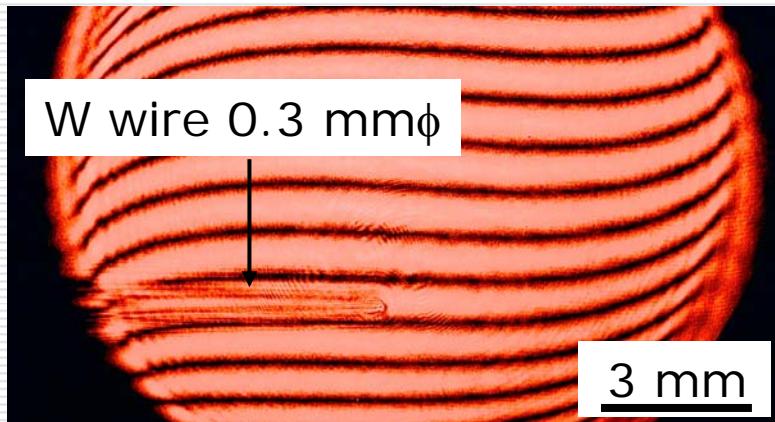
Refractive index: 1.00002714

Heating

Temperature: 323 K (50°C)

Gas: Ar 1×10^4 Pa

Refractive index: 1.00002503



We can detect only a difference of 10^{-6} - 10^{-7} orders!!

$$n_{Ar(T,P)} - 1 = \frac{[n_{Ar(273.15, P_0)} - 1]}{1 + a\Delta T} \frac{P}{P_0}$$

$$\Delta d = \{ (n(T_0, P_0) - n(T, P)) \} \times l / \lambda .$$

a : the coefficient of volume expansion,

Δd : degree of movement,

P_0 : 101333.25 Pa

l : the physical length

In-situ observation of nucleation: WO_3

Temperature: 298 K (25°C)

Gas: Ar 1×10^4 Pa

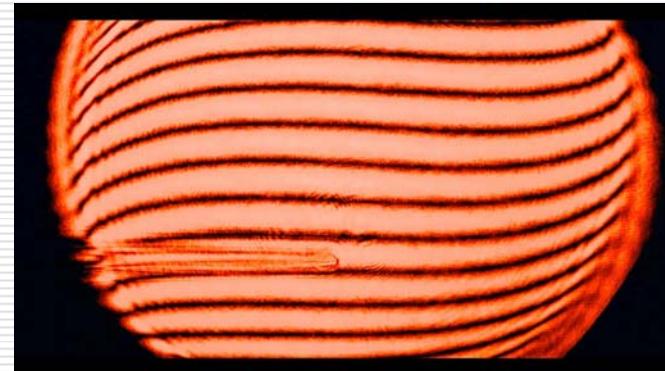
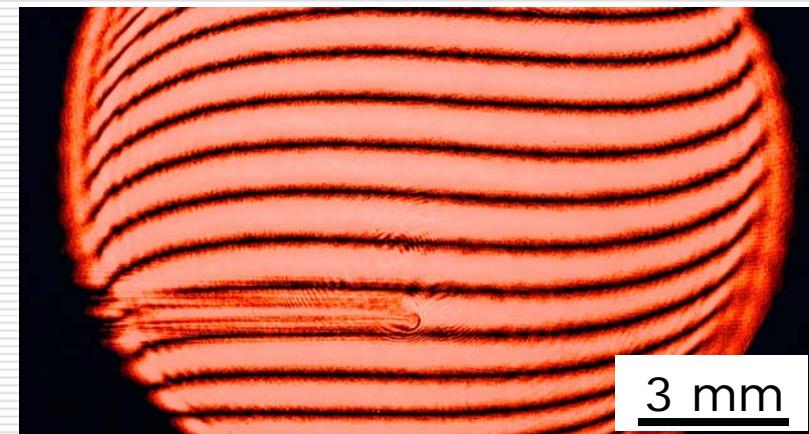
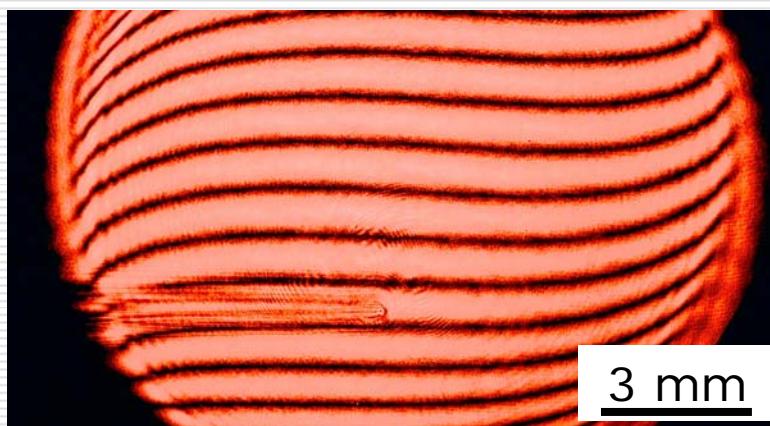
Refractive index: 1.00002714

Oxygen

Temperature : 298 K (25°C)

Gas : Ar 9×10^3 Pa, O₂ 1×10^3 Pa

Refractive index: 1.00002703



Only Temperature

Heating
RT
↓
1570 K



Temperature & concentration

Temperature information is subtracted by oxygen free experiment.

In-situ observation of nucleation: Mn



5 times fast

in Ar: 2.0×10^4 Pa

Evaporation Source

T: 1953 K

P_e: 1.6×10^4 Pa

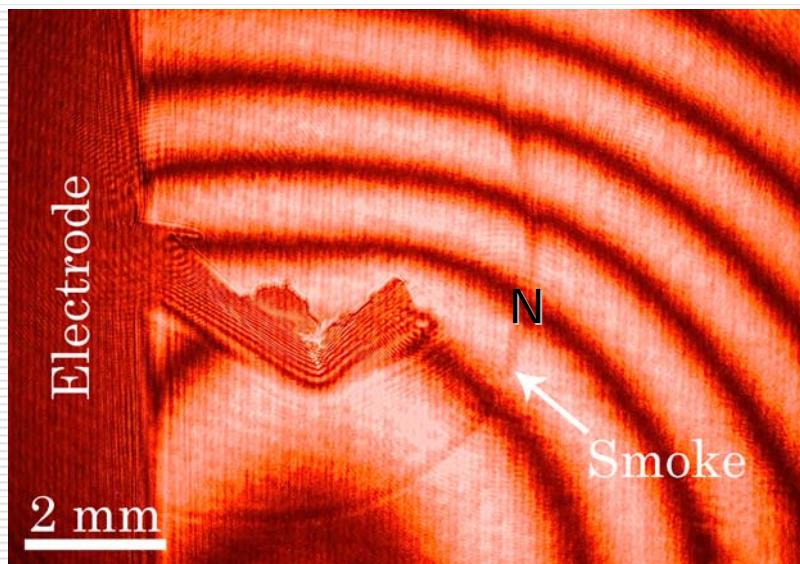
Nucleation site "N"

T: **1100 K**

P: 2.7×10^3 Pa

P_e: 4.4×10^{-2} Pa

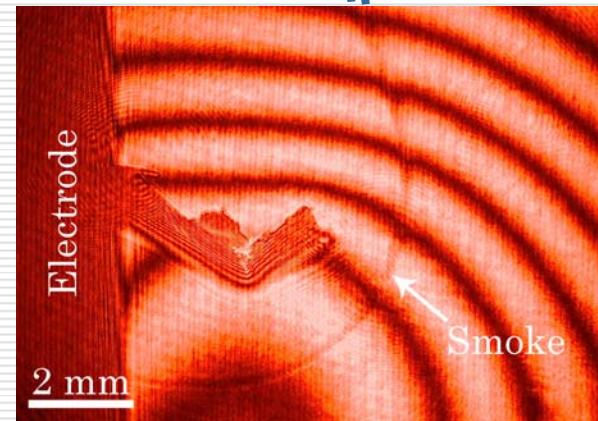
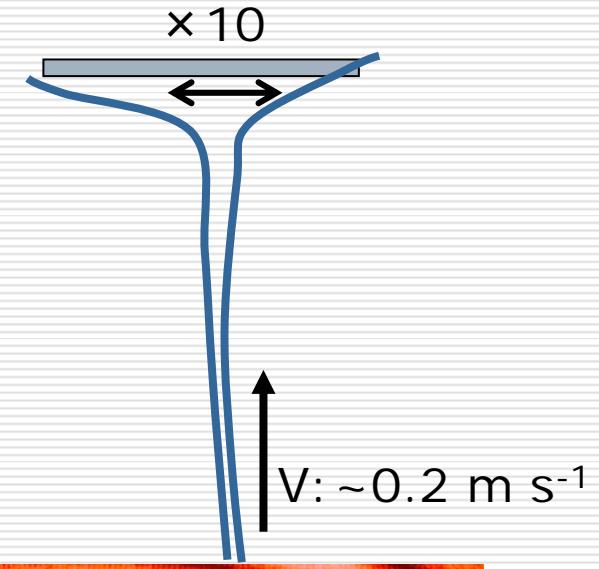
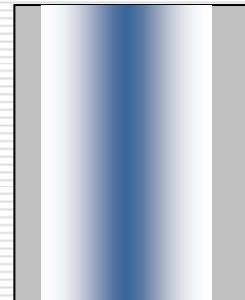
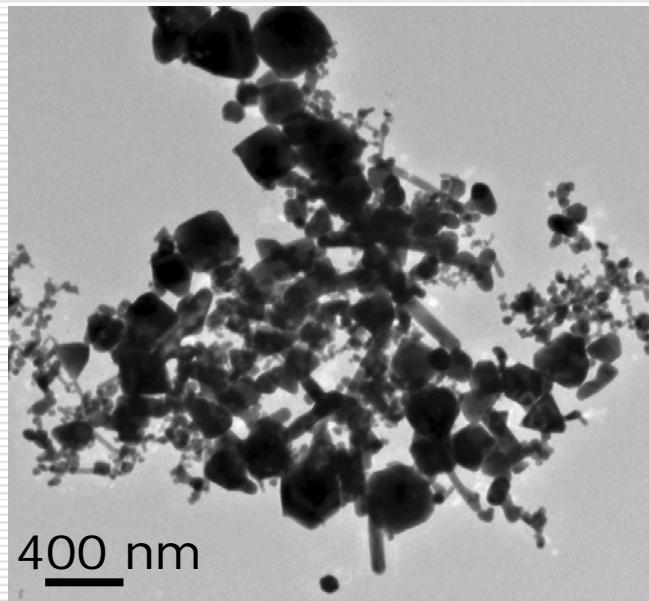
$$P/P_e = \mathbf{6.1 \times 10^4}$$



Condensation temperature (1100 K) is significantly lower than their equilibrium temperature (1760 K) at 2.7×10^3 Pa.

$$\Delta T = \mathbf{660 \text{ K}}$$

In-situ observation of nucleation: Mn



Mean radius: 13.7 nm

Number density: $6 \times 10^{15} \text{ m}^{-3}$

$$= 10 \times \rho / (V \times t)$$

ρ : Number density on the grid (m^{-2})

V : Velocity of smoke (m s^{-1})

t : Experimental duration (s)

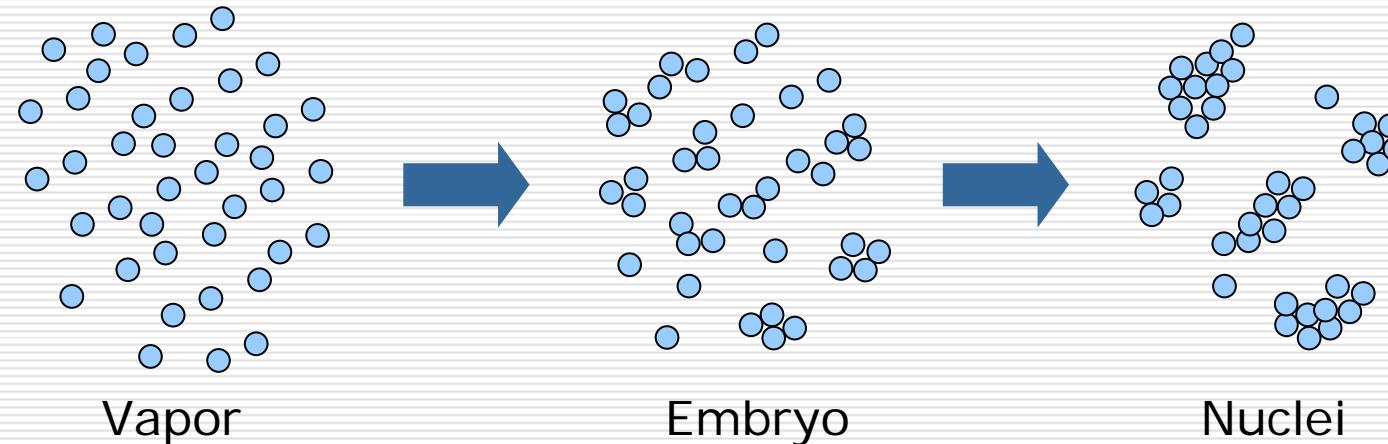
Is there useful nucleation theory?

Classical nucleation theory does not work in nano-scale.

Indeed, nucleation rate from classical theory is far from experimental & MD simulation values.

(water, ethanol, etc) (water, inert gas)

(Feder 1966, Schmitt et al. 1982,
Adams et al. 1984, Yasuoka et al.
1998, Toxvaerd 2001, 2003)



Must consider intermolecular interaction.

Is there useful nucleation theory?

Semi-phenomenological nucleation theory: SP

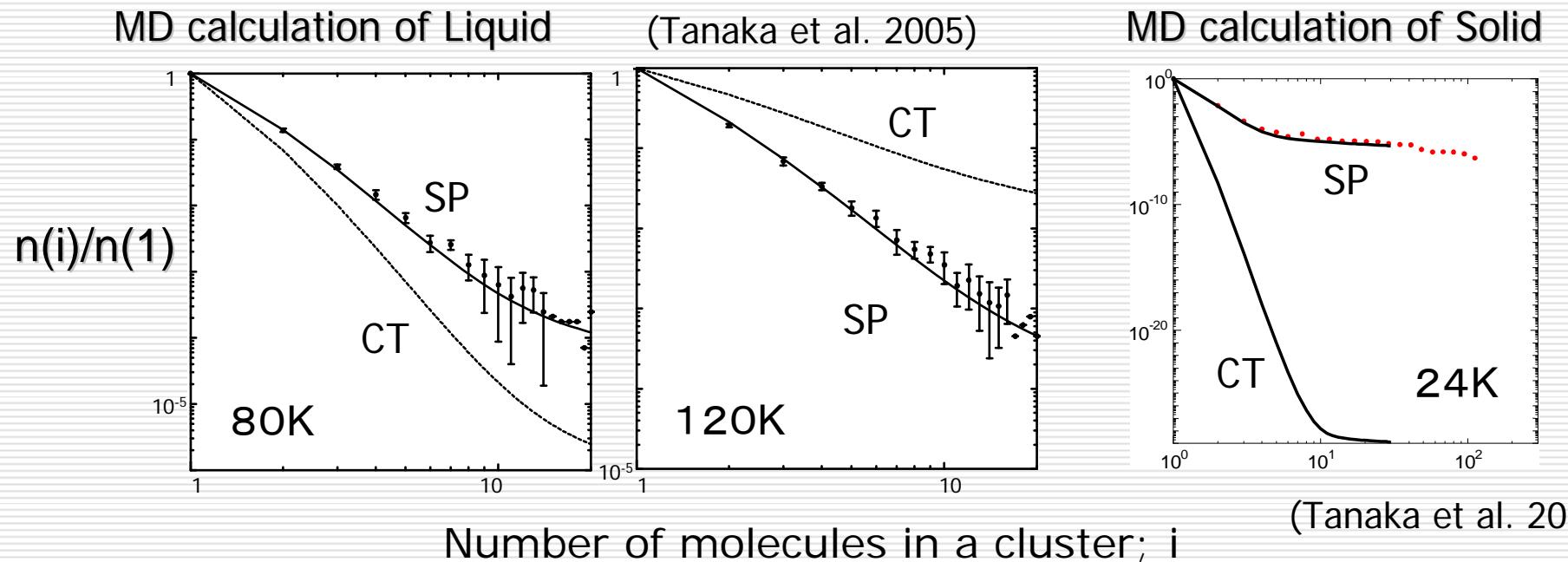
(Dillmann & Meier 1991)

$$\mu_i = kT \log(P_i/P_e) + i\mu_{\text{solid}} + \sigma A_1 i^{2/3} + \underline{A_2 i^{1/3}} - \underline{\sigma A_1 \cdot A_2}$$

Correction term

Correction term: Determined by second virial coefficient.

σ : Surface energy per unit area



Values for model calculation

- Initial pressure of Mn: $P_0 = 2.7 \times 10^3$ Pa
- Time scale for cooling: $\tau = 1.12 \times 10^{-3}$ s

Experiment

- Surface free energy of molten Mn at 1573-1773K: 1100 erg/cm²

Physical Chemistry of High Temperature Technology, E.T. Turkdogan, Academic Press, New York, 1980

- Dissociation energy of dimer: 2700 K

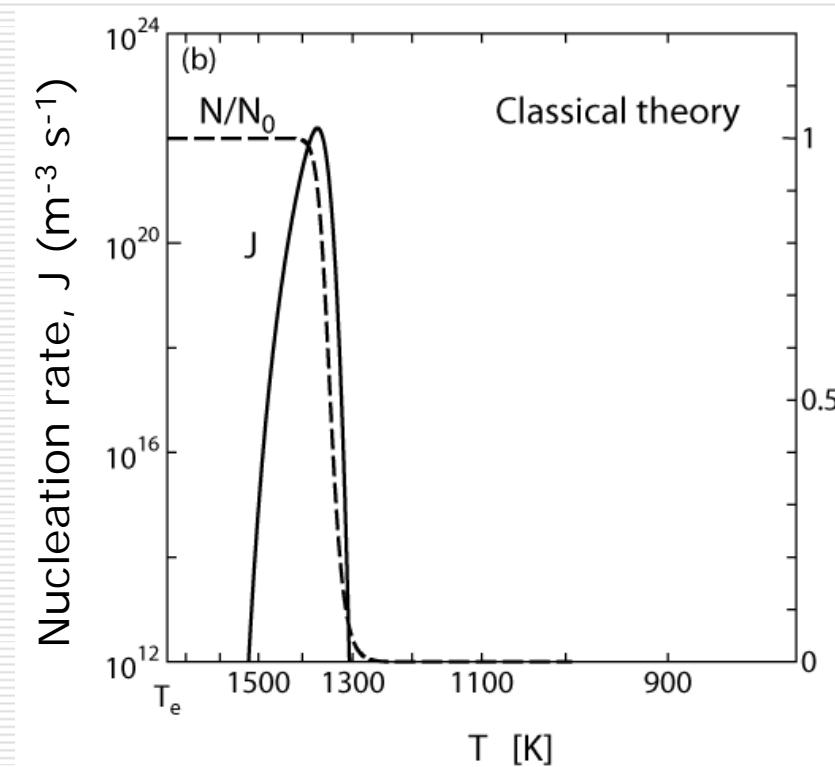
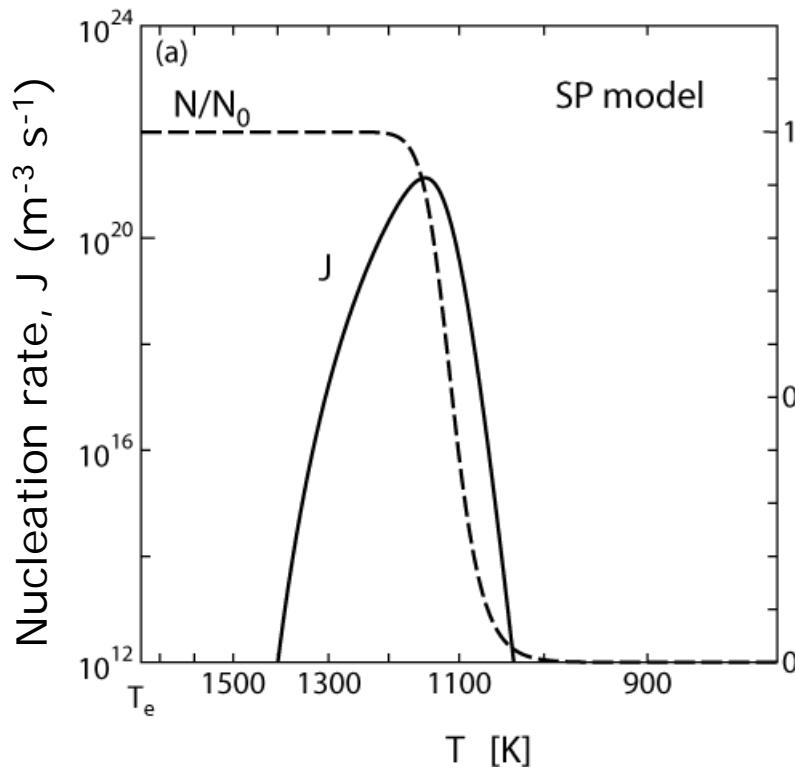
A. Kant, S.-S. Lin, and B. Strauss
Dissociation energy of Mn₂ J. Chem. Phys. vol.49, p1983-1985, 1968

- Equilibrium distance between nuclei: $R_e = 2.34 \times 10^{-8}$ cm

- Vibrational frequency: 1×10^{13} s⁻¹

Dissociation energies of diatomic molecules of the transition elements. II Titanium, Chromium, Manganese, and Cobalt J. Chem. Phys. vol.41, p.3806-3808, 1964

Condensation Temperature



Experiment

Condensation T : 1100 K

Number density: $6 \times 10^{15} \text{ m}^{-3}$

Mean radius : 13.7 nm

SP Model

1110 K

$7.5 \times 10^{16} \text{ m}^{-3}$

16 nm

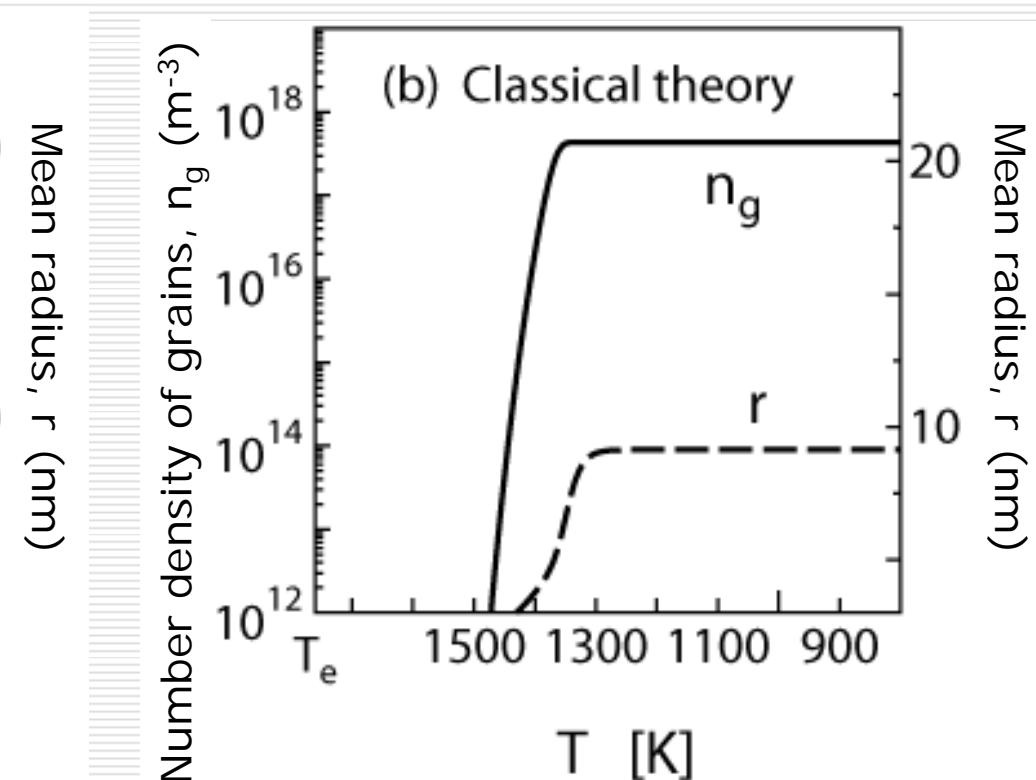
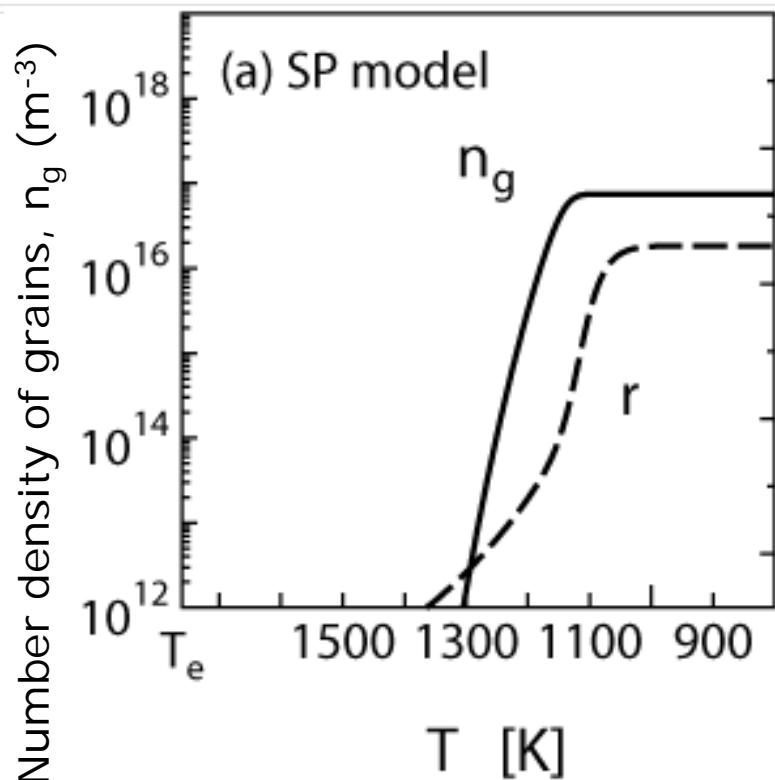
CNT

1340 K

$4.3 \times 10^{17} \text{ m}^{-3}$

9.1 nm

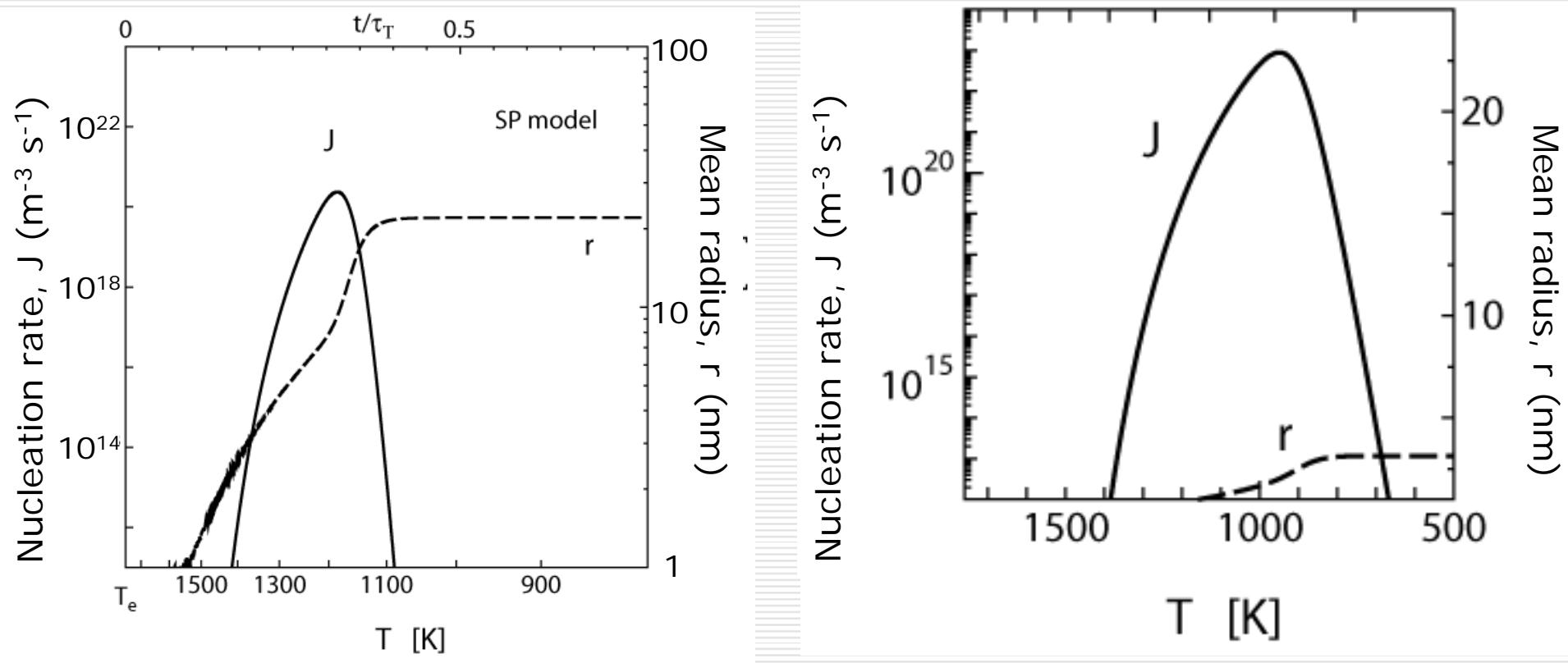
Number density & Size



Experiment	SP Model	CNT
Condensation T : 1100 K	1110 K	1340 K
Number density: $6 \times 10^{15} \text{ m}^{-3}$	$7.5 \times 10^{16} \text{ m}^{-3}$	$4.3 \times 10^{17} \text{ m}^{-3}$
Mean radius : 13.7 nm	16 nm	9.1 nm

Surface free E can be determined using condensation T .

Validity of a Sticking coefficient



	Experiment	SP Model	SP Model
Sticking probability:		1.0	0.1
Condensation T :	1100 K	1110 K	898 K
Mean radius :	13.7 nm	16 nm	3.1 nm

Validity of a Sticking coefficient

We assumed the sticking coefficient is unity, which will be confirmed from growth rate and timescale for growth.

▪ **Growth velocity:** $R = \frac{\alpha \Omega (P - P_e)}{(2 \pi m k T)^{1/2}}$, Hertz-Knudsen's equation.

Ω (volume of molecule): $1.22 \times 10^{-23} \text{ cm}^{-3}$

P (vapor pressure): $2.7 \times 10^4 \text{ dyne cm}^{-2}$

P_e (equilibrium vapor pressure): $4.4 \times 10^{-1} \text{ dyne cm}^{-2}$

m (molecular weight): $9.11 \times 10^{-23} \text{ g}$

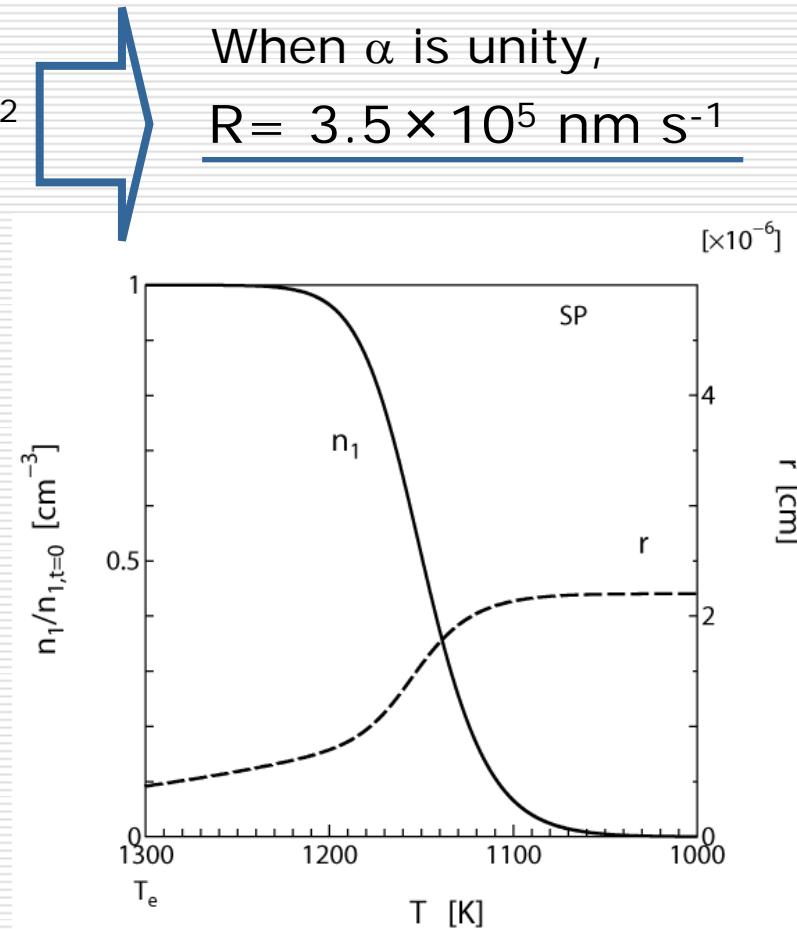
K (Boltzmann constant): $1.38 \times 10^{-16} \text{ erg K}^{-1}$

T (absolute temperature): 1100 K

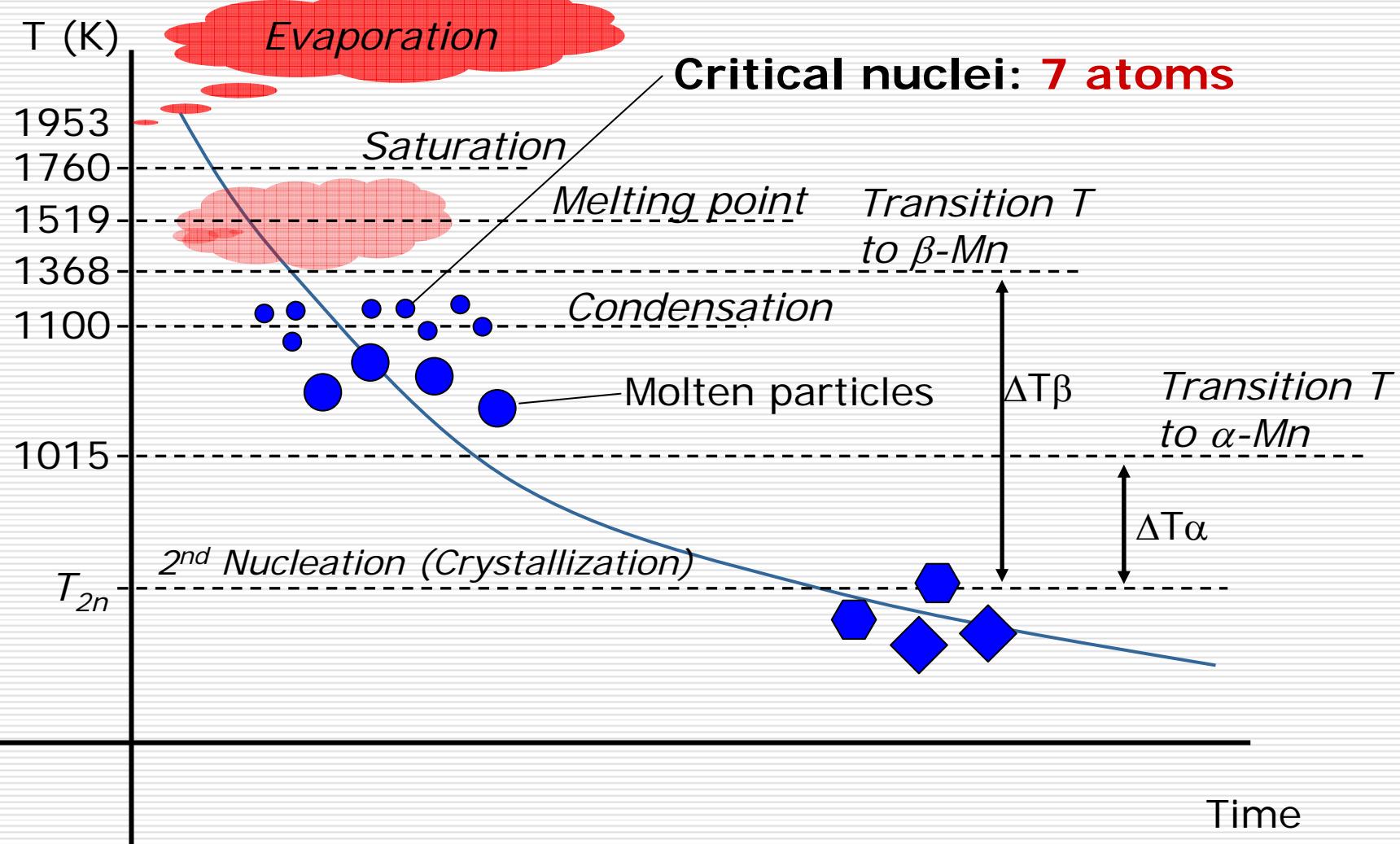
Time scale for grains growth corresponds to timescale for gas cooling for 100 K, $1 \times 10^{-4} \text{ s}$.

Mn particles can be grown to 35 nm.

α will be $14 / 35 = 0.4$.



Formation process of Mn nanoparticles



Conclusion

- Condensation occurs under very high supercooling ($\Delta T=660K$).
- Supersaturation is as high as 10^4 - 10^5 .
- SP model works well.
- Sticking probability can be determined.

In Dec., 2011



Dec., 2012

