

# Why evolution?

Why the Stars and the Universe Evolve?

@Kobe planetary school

Daiichiro Sugimoto

former affiliations: Univ. Tokyo

& Open University of Japan

Jan 10, 2011

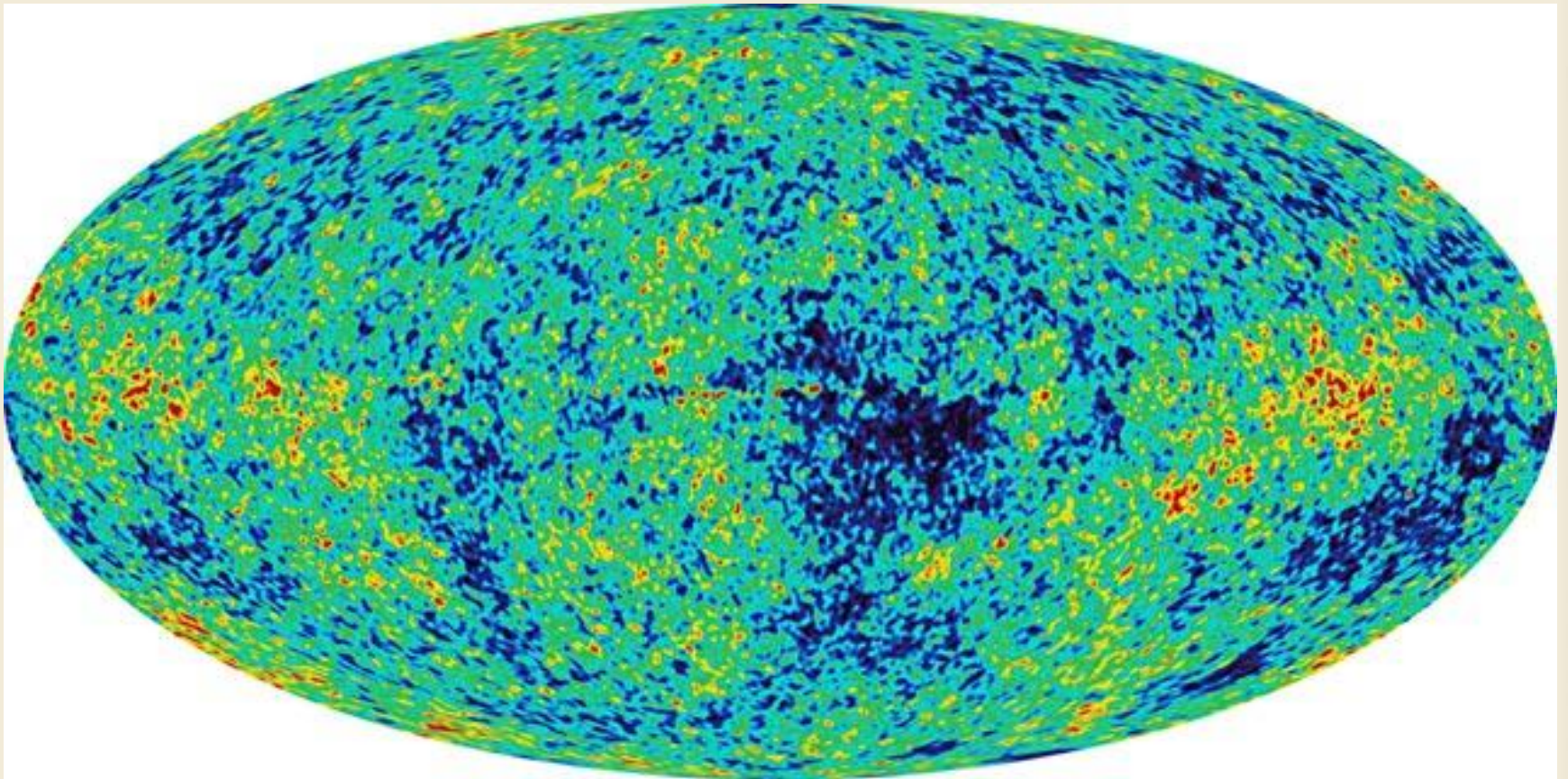
# Wärmetod or Emergence of structures?

- R. Clausius: Thermal death of the universe (1865)
- **In reality:** Full of diversities (structures)
- Initially non-(thermal) equilibrium ?
  - Who gave such information initially? God?
- Contrary, 3K bkgd radiation tells initially almost in eql:
  - extent of its fluctuation is **only** 1/100000
  - **requires emergence of orders (information)**
- What are the mechanisms to **generate** structures and information of such a great extent (contrast)
- **There are two important mechanisms**

# Wilkinson Microwave Anisotropy Probe

WMAP : by NASA

$10^{-5}$



# The first mechanism

- Boundary Conditions to the system change
  - Universe expanded more rapidly than recovering therm. eql.
    - **time scales: expansion** << **relaxation** in early Univ
  - NSE composition: H (high T) / He / Fe (low T)
    - temp dropped before the nuclear eql re-established  
i.e., **dropping-out** (falling-behind) of changing thermal eql
- Other examples:
  - earth's atmosphere out of therm eql
    - ascending current is rapid
  - (in industry) fractional distillation of crude oil

# The second mechanism

-- effects of **gravity** (**gravothermodyn**)

- Forming core-halo structures
  - in density and/or temperature distributions:  
**celestial obj / thin space**, clustering of obj's,  
stellar core / halo, **the sun / the planets**
  - in specific angular momentum distr:  
galactic core / spiral arm, **the sun / the planetary sys**
- Mechanism: Coupling between gravity and thermal effects (kinetic energy – pressure)

## Simple example: Virial Theorem

$$\frac{dP}{dr} = -\frac{GM_r \rho}{r^2} \quad \Rightarrow \quad \int_0^R \frac{dP}{dr} 4\pi r^3 dr = - \int_0^R \frac{GM_r \rho}{r^2} 4\pi r^3 dr$$

partial integr & BC with  $P = (\gamma - 1)\rho u$

$$3(\gamma - 1) \int_0^M u dM_r - \int_0^M \frac{GM_r}{r} dM_r = 0$$

$$2E_{\text{th}} + E_{\text{gr}} = 0 \quad \text{for } (\gamma = 5/3)$$

$$E_{\text{tot}} = E_{\text{gr}} + E_{\text{th}} = -E_{\text{th}} < 0$$

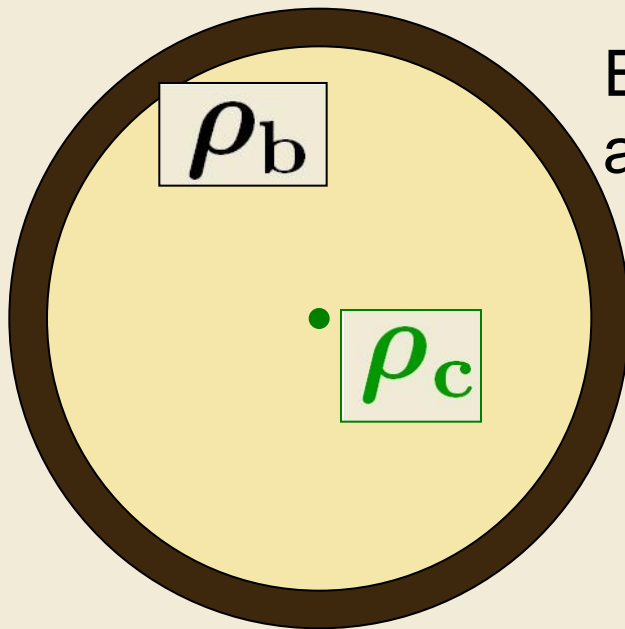
$$L = -\frac{dE_{\text{tot}}}{dt} = \frac{dE_{\text{th}}}{dt} > 0$$

Temperature rises  
when energy (entropy) dumped away  
- negative (effective) specific heat

# Gravothermal catastrophe

- В.А. Антонов, [1962](#), Вест. Ленингр. Госдр. Унив.
- D. Lynden-Bell & R. Wood, [1968](#), MN (turn pt of lin series)
- I.Hashisu & D.Sugimoto, [1978](#), PTP  
normal-mode method
- I.H., Y.Nakada, K.Nomoto & D.S., [1978](#), PTP  
catastrophe of finite amplitude = grav. contr.
- D.S., Y.Eriguchi & I.H., [1981](#), PTP Suppl.  
summary and extension to cosmology
  - easy understandings of stellar structure & evolution
  - resolved the problem of thermal death of the univ

# gravothermal catastrophe (isothermal gas sphere)



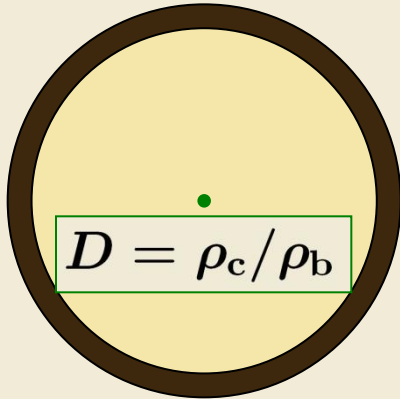
Enclosed by  
adiabatic wall

density contrast

$$D \equiv \frac{\rho_c}{\rho_b} > 1$$



# Redistribution of specific entropy



Removal of heat (specific entropy)

$$\delta q = T\delta s < 0$$

$$\delta T = \frac{\delta q}{c} + (\delta T)_{\text{comp/exp}} = \frac{\delta q}{c^*}$$

Change due to resulting compression

$$(\delta T)_{\text{comp/exp}} > 0$$

Combined result can be  $\frac{\delta q}{c^*} < 0$

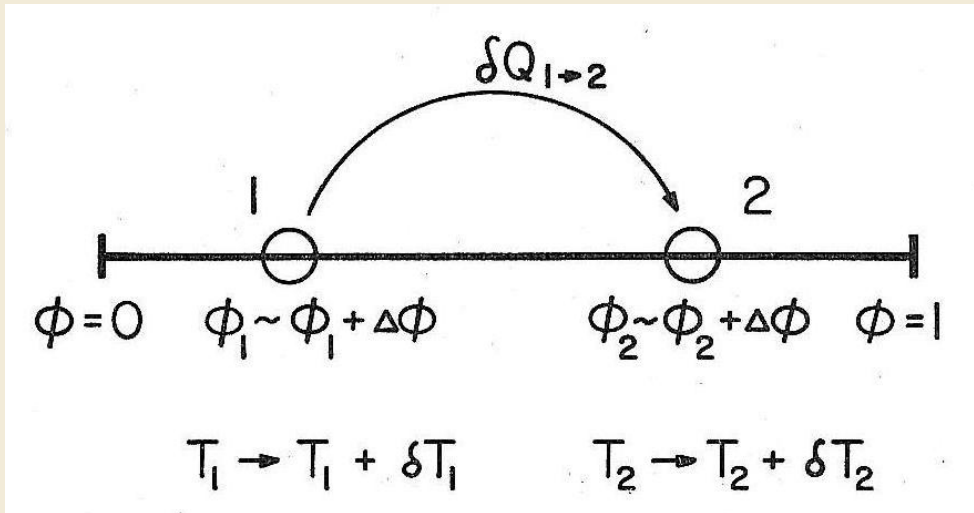
$c^*$  : gravothermal specific heat

- When  $D < 709$  :  $c^* > 0$ , i.e., thermal state
- When  $D > 709$  :  $c^* < 0$ , i.e., gravothermal state

↑ negative specific heat

leading to thermodynamic instability

# Negative specific heat leads to thermal instability



first order:

$$\delta S = \left( \frac{1}{T_2} - \frac{1}{T_1} \right) \delta Q_{1 \rightarrow 2} = 0$$

for initially isothermal case

second order:

$$\delta^2 S = \left( \frac{1}{T_0 + \delta T_2 / 2} - \frac{1}{T_0 + \delta T_1 / 2} \right) \delta Q_{1 \rightarrow 2} = \frac{1}{2T_0} \left( \frac{\delta T_1}{T_0} - \frac{\delta T_2}{T_0} \right) \delta Q_{1 \rightarrow 2} > 0$$

for  $\delta T_1 > \delta T_2$  (for negative specific heat)

i.e., entropy production (unstable)

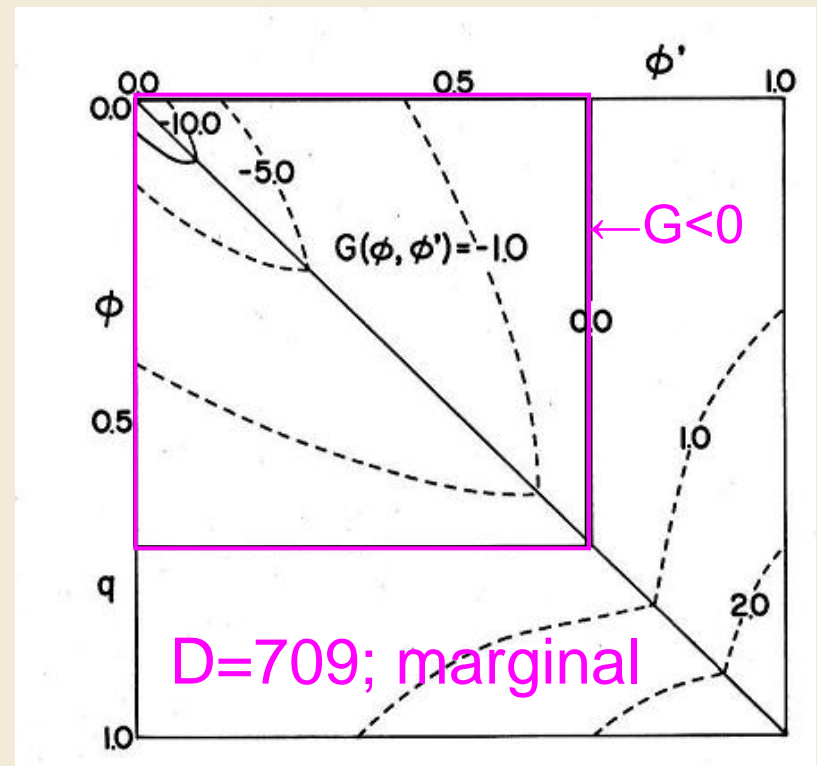
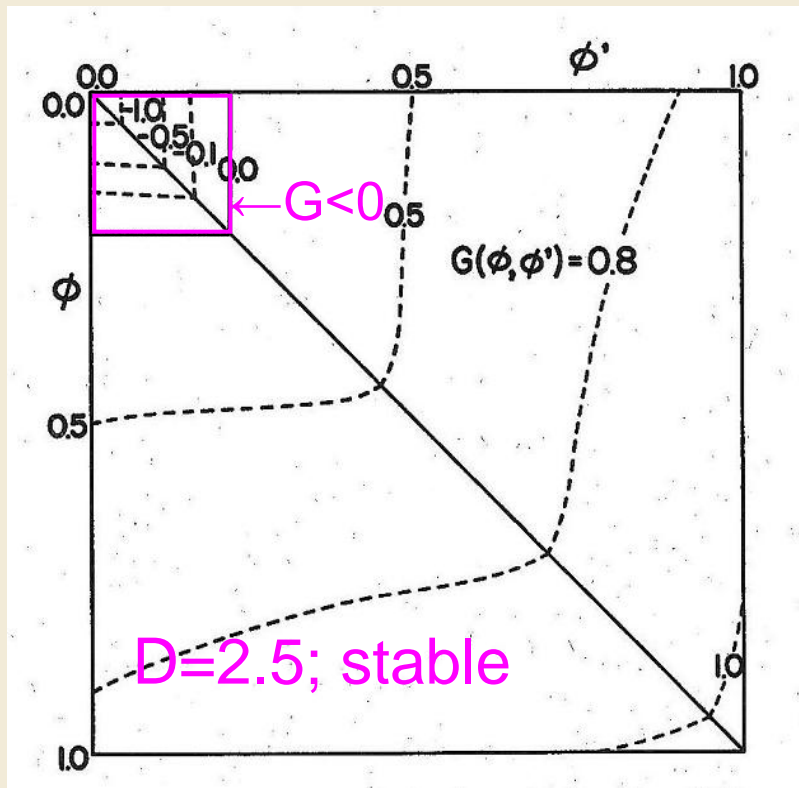
→ Temp gradient spontaneously appears

inverse  
(gravothermal)  
tensor  
specific heat

$$\delta \ln P = \int_0^1 G(\phi, \phi') \delta \sigma(\phi') d\phi'$$

$$\delta \ln T = \int_0^1 F(\phi, \phi') \delta \sigma(\phi') d\phi'$$

$$F(\phi, \phi') = \frac{2}{5} [G(\phi, \phi') + \delta(\phi - \phi')]$$



## Fundamental reason

phenomena:	energy	common sense of thermodynamics?
------------	--------	------------------------------------

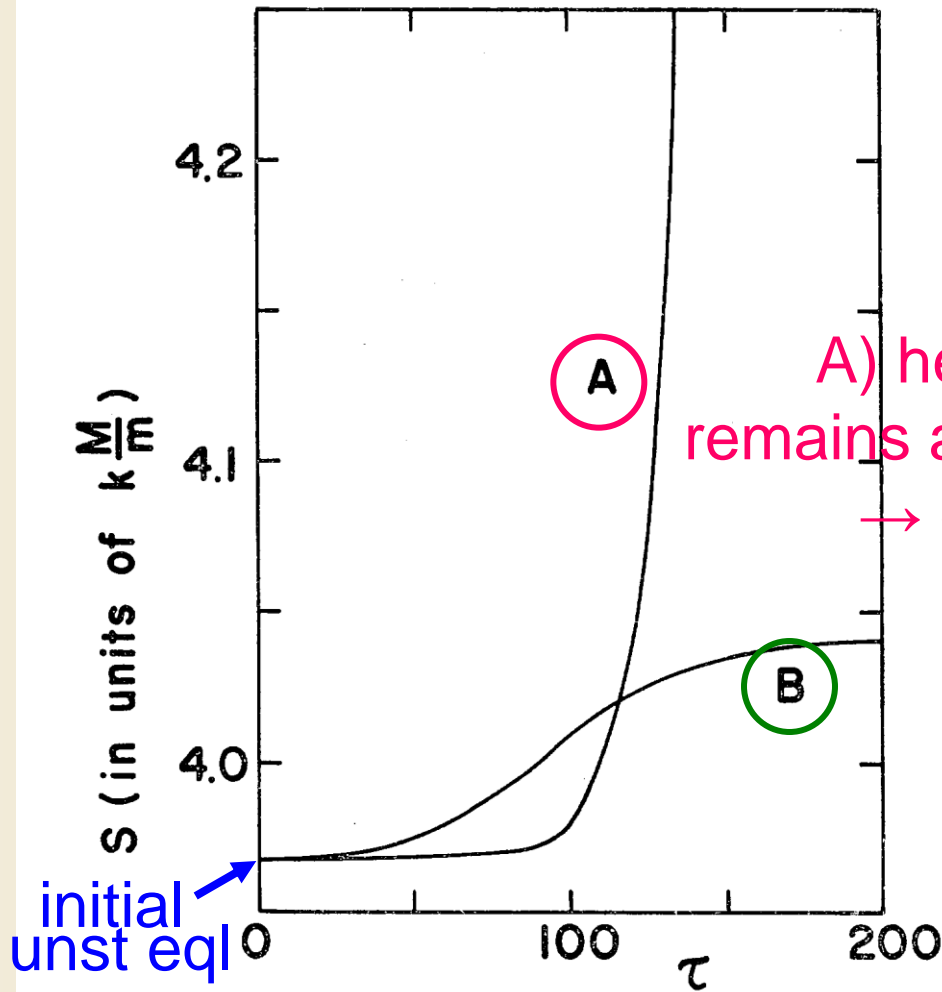
---

▪ surface tension:	$M^{2/3}$ (sub-extensive)	along
--------------------	------------------------------	-------

▪ general cases:	$M^1$ (extensive)	along
------------------	----------------------	-------

▪ gravitational:	$M^2$ (super-extensive)	out of
------------------	----------------------------	--------

coming fm: self energy  $\leftarrow \infty$  effective range



Two cases A) & B)  
of evolution for different  
initial perturbations

A) heat outwards  
remains as gravoth system  
→ catastrophe

B) heat inwards  
towards thermal system  
→ another stable eql

Fig.5. Changes in the total entropy of the system for the cases A and B.

## Fate of Evolution

Initially in Gravo-thermal state  
evolution triggered by

(A) outward heat flow

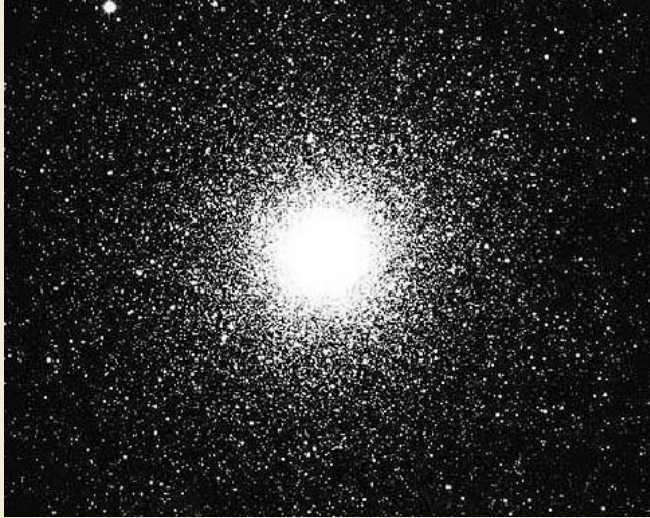
- core contr vs halo exp
- T and  $\rho$  gradients more enhanced
- higher D → more gravo-thermal
- $\rho_c \rightarrow \infty$  in a finite time, (catastrophe)

(B) inward heat flow

- core exp & halo contr
- lower D → less gravo-thermal
- isothermal state

(another thermal state with higher S )

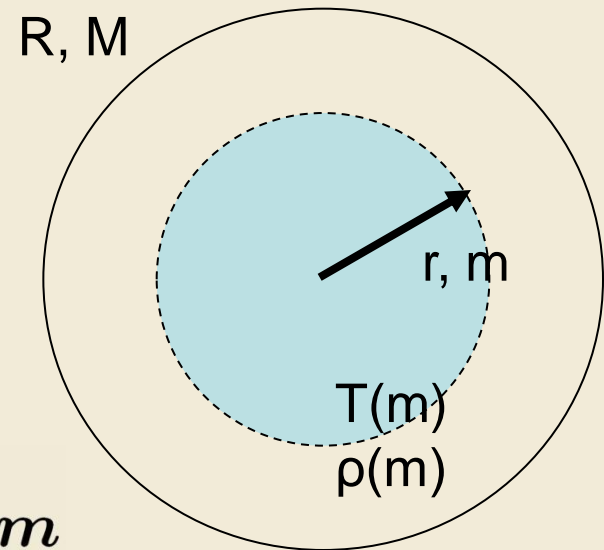
## Applied for open system



- A globular star cluster  
Gravitational N-(many-) body problem:  
Each star represented by a point-mass  
gravitationally interacting each other.

- Stellar Interior  
T &  $\rho$  distributions  
as fns of Lagrangian co-ordinate  $m$

In both cases: 
$$S = \int_0^M s(m) dm$$



## Entropy Aspect of the System

$$\frac{dS}{dt} = \left(\frac{dS}{dt}\right)_{\text{in}} - \left(\frac{dS}{dt}\right)_{\text{out}} + \left(\frac{dS}{dt}\right)_{\text{irr}}$$

Open system (Stellar interior / Earth atmos / Life =metabolism)

*input* fm ( Ln / solar rad / food )

*out* to ( stellar light / IR rad / heat & elimination)

*irr.* entropy production by heat transport etc

- (secular) cooling (grav contraction): no input & S decreasing

- **Steady state**  $\frac{dS}{dt} = 0$  (nuclear burning phase)

- (rapid) thermodyn transition

$\left(\frac{dS}{dt}\right)_{\text{out}}$  small as comp'd with other terms (flash etc)



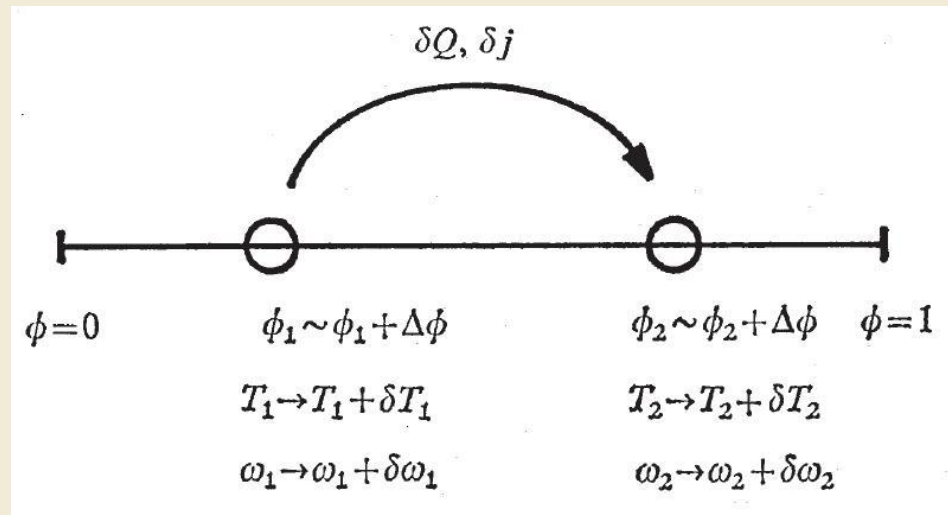
# Gravothermal system with rotation

## Parallelism

intensive variable	T	$\omega$
extensive variable	s	j

W: Liapunov fn

$$\delta^2 W = - \int_V (\delta T \delta s + \delta \omega_z \delta j_z) da < 0 ,$$



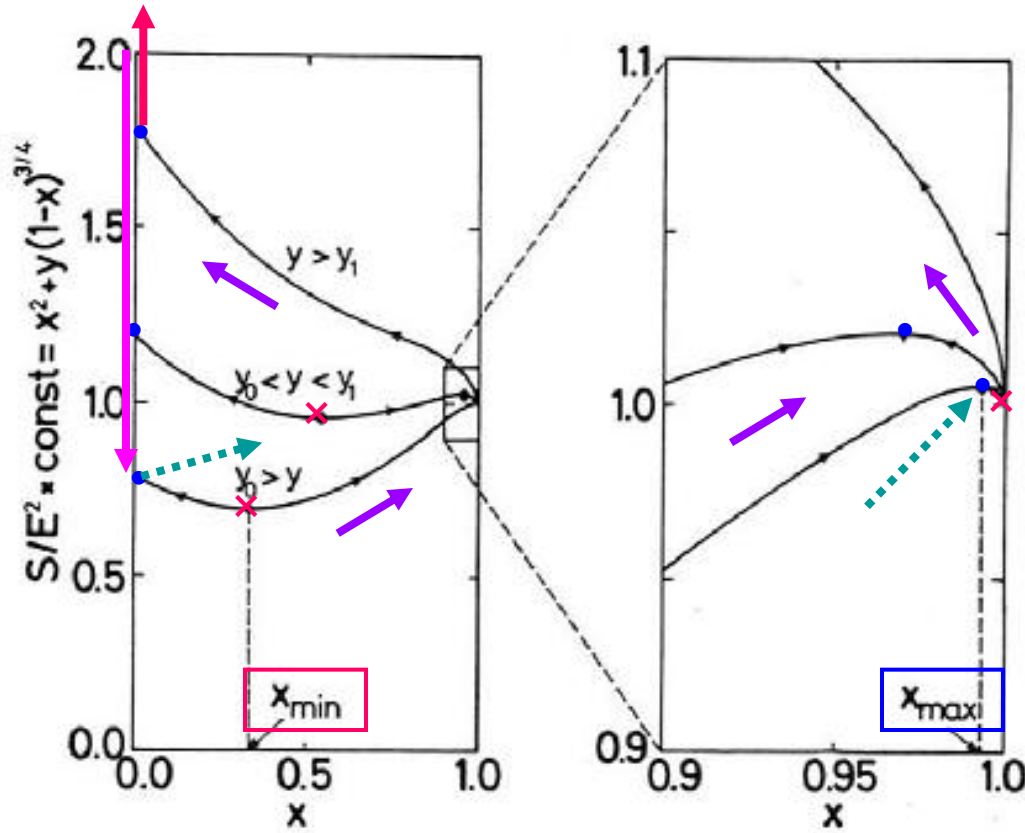
Initial equilibrium state holds by balance among grav force, pressure grad, & centrifugal force

initially:  
 $T = \text{const}$   $\omega = \text{const}$   
↓  
gradient appears spontaneously :  
further proceeds  
 $s\text{-flow}$   $j\text{-flow}$   
↓  
the gradient grows unstably

similar to artificial earth satellite  
losing its kinetic energy in atmosphere  
losing its  $j$  (transferring outwards)  
→ its total energy increases,  $\omega$  increases  
gradient in  $\omega$  increases

# Cosmology

Sugimoto et al, Prog Theor Phys Suppl #70, 1981)



## BH + rad

- (local) max
- × (local) min

within particle horizon

- inflation (decrease in E)
- early expansion (increase in E)

for comoving volume

- rapid phase transition (E=const; nucleus formed by gravotherm)

$$x \equiv M_{\text{BH}} c^2 / E,$$

$$E \equiv M_{\text{BH}} c^2 + a T_r^4 V$$

$$S_{\text{tot}} = S_{\text{BH}} + S_r$$

$$y \equiv [a (T_g / 3\pi)^4 V / E]^{1/4} (m_g c^2 / E)$$

# Our World

## Nesting of non-equilibria

### Non-eql universe due to expansion

- ▷ Nuclides out of equilibrium, stars vs interstellar space driving stellar evolution
  - ▷ The sun and the earth
    - ▷ Different Temperatures between earth surf and top of atmosphere
      - ▷ Atmosphere of unsaturated humidity
        - ▷ Lives as non-equilibrium open system in steady state

# Reference 1

- I.Hashisu,Y.Nakada,K.Nomoto & D.Sugimoto, 1978;Gravothermal Catastrophe of Finite Amplitude,PThPh,60,393
- D.Sugimoto,Eriguchi Y.,Hachisu I.,1981,Gravothermal Aspects in Evolution of Stars and the Universe,PThPh,70,154
- <http://map.gsfc.nasa.gov/>
- D. Lynden-Bell & R. Wood, 1968; The gravo-thermal catastrophe in isothermal spheres and the onset of red-giant structure for stellar systems,MNRAS,138,495